

AperTO - Archivio Istituzionale Open Access dell'Università di Torino

Quantum Mechanical Investigations on the Formation of Complex Organic Molecules on Interstellar Ice Mantles. Review and Perspectives

This is the author's manuscript

Original Citation:

Availability:

This version is available <http://hdl.handle.net/2318/1728937> since 2020-02-19T21:04:50Z

Published version:

DOI:10.1021/acsearthspacechem.9b00082

Terms of use:

Open Access

Anyone can freely access the full text of works made available as "Open Access". Works made available under a Creative Commons license can be used according to the terms and conditions of said license. Use of all other works requires consent of the right holder (author or publisher) if not exempted from copyright protection by the applicable law.

(Article begins on next page)

This document is confidential and is proprietary to the American Chemical Society and its authors. Do not copy or disclose without written permission. If you have received this item in error, notify the sender and delete all copies.

**Quantum Mechanical Investigations on the Formation of
Complex Organic Molecules on Interstellar Ice Mantles.
Review and Perspectives**

Journal:	<i>ACS Earth and Space Chemistry</i>
Manuscript ID	sp-2019-000827
Manuscript Type:	Review
Date Submitted by the Author:	01-Apr-2019
Complete List of Authors:	Rimola, Albert; Universitat Autònoma de Barcelona, Química Zamirri, Lorenzo; Università degli Studi di Torino Dipartimento di Chimica, Chemistry Ugliengo, Piero; Università degli Studi di Torino, Dipartimento di Chimica Ceccarelli, Cecilia; Univ. Grenoble Alpes, CNRS, Institut de Planétologie et d'Astrophysique de Grenoble (IPAG)

SCHOLARONE™
Manuscripts

Quantum Mechanical Investigations on the Formation of Complex Organic Molecules on Interstellar Ice Mantles. Review and Perspectives

Lorenzo Zamirri,^{1,2} Piero Ugliengo,^{1,2} Cecilia Ceccarelli,³ and Albert Rimola.^{4*}

¹*Dipartimento di Chimica, Università degli Studi di Torino, via P. Giuria 7, 10125, Torino, Italy*

²*Nanostructured Interfaces and Surfaces (NIS) Centre, Università degli Studi di Torino, via P. Giuria 7, 10125, Torino, Italy*

³*Univiversité Grenoble Alpes, CNRS, Institut de Planétologie et d'Astrophysique de Grenoble (IPAG), rue de la Piscine 414, 38000, Grenoble, France*

⁴*Departament de Química, Universitat Autònoma de Barcelona, 08193, Bellaterra, Catalonia, Spain*

*Corresponding author: albert.rimola@uab.cat

Abstract

The interstellar medium (ISM) is rich in molecules, from simple diatomic to complex organic ones, some of which have a biotic potential. A notable example, in this respect, is represented by the so-called interstellar complex organic molecules (iCOMs). Interestingly, the various phases involved in the formation of Solar-type planetary systems lead to an increasing chemical complexity, in which, at each step, more complex molecules form. In dark molecular clouds, dust grains are covered by ice mantles, mainly made up of H₂O but also of other volatiles species such as CO, NH₃, CO₂, CH₄ and CH₃OH. Although their mass is one hundred times lower than the gas-phase matter, these ice-covered grains play a fundamental role in the interstellar chemical complexity as some important reactions are exclusively catalyzed by their surfaces. For example, one of the current paradigms on the iCOMs formation assumes that iCOMs are synthesized on the ice mantle surfaces, in which reactants accrete and diffuse to finally react. As the usual approaches employed in astrochemistry (*i.e.*, spectroscopic astronomical observations, astrochemical modelling and laboratory experiments) cannot easily provide details on the iCOMs formation processes occurring on ice mantles at the atomic level, computational chemistry has recently become a complementary tool to fill in this gap. Indeed, it can provide an accurate description (*i.e.*, structures and reactive

energy profiles) of these processes. Accordingly, several recent studies simulating the formation of iCOMs on icy surfaces by means of quantum mechanical methods have appeared in the literature. This article aims to comprehensively review most of these works, focusing not only on standard iCOMs but also on simpler organic compounds as well as biomolecules. Perspectives on possible future directions of research using computational chemistry are also proposed.

1 Introduction

Despite the harsh conditions of the interstellar medium (ISM), more than 200 interstellar molecules have been discovered so far, with this number steadily increasing with time.¹ Among them, the class of C-bearing molecules with at least 6 atoms are defined as interstellar complex organic molecules (iCOMs).^{2,3} Such a definition allows us to exclude simple molecules that are sure not *organic*, like H₂O, NH₃ or CO, but excludes some relevant ones like formaldehyde (H₂CO) and methanimine (CH₂=NH). Moreover, other species, although not being categorized as iCOMs, can play a crucial role in the organic, and eventually pre-biotic, chemistry occurring in the ISM, such as the case of formic acid (HCOOH), hydrogen cyanide/isocyanide (HCN/HNC) or the isocyanic acid (HCNO), just to mention a few.

Nonetheless, iCOMs have lately received a lot of attention for their potential contribution to the emergence of life and because iCOMs in solar-type hot corinos provide a direct link between interstellar chemistry and the small bodies of the Solar System, *i.e.* comets and asteroids.⁴⁻⁶

Although the presence of iCOMs has been known for decades,⁷ the chemical routes that lead to their formation is still matter of intense debate. Two alternative paradigms are invoked in the literature: either iCOMs form in the gas-phase,⁸⁻¹⁰ or on the interstellar grain surfaces.^{8,11-13} Here, we will focus on the latter paradigm. Briefly, it postulates that iCOMs are synthesized on the grain surfaces following a three-step process: *i*) hydrogenation of frozen atoms and molecules to form saturated species (*e.g.*, CH₃OH from CO^{14,15}) during the cold prestellar phase; *ii*) formation of radicals (*e.g.*, CH₃O·, HCO·, NH₂·) derived from the frozen hydrogenated species due to incidence of UV radiation and cosmic rays on the ice mantles, and *iii*) coupling of radicals to form iCOMs, in

1
2
3 which radicals are assumed to diffuse on the ice mantles due to temperature increase (≈ 30 K) during
4 the protostellar phase.^{11–13}
5
6

7
8 In this paradigm, therefore, interstellar grains play a major role, so we briefly describe their
9 characteristics, in the context of this review. They are silicate and carbonaceous sub-micron size
10 particles that permeate most of the Galaxy ISM.^{16–18} In cold molecular clouds, grains are enveloped
11 by iced mantles constituted mostly of water with smaller quantities of carbon monoxide and dioxide
12 (CO, CO₂), ammonia (NH₃) and methanol (CH₃OH).^{19,20} Therefore, when referring to grain surface
13 reactions, what one really means is reactions occurring on surfaces of water ice.
14
15
16
17
18
19
20
21

22 Little is known about the structure of these iced mantles. Experiments evidenced similarities
23 between the IR features of interstellar ices and those of amorphous solid water (ASW),²¹ and
24 accordingly they are usually referred to be amorphous and partly porous.²² However, IR
25 spectroscopy is not a definitive technique to derive conclusive structural features of these ices, as
26 outlined in a recent work on the solid CO/H₂O interface.²³ From IR observations, it seems that ices
27 present two different solid phases: *i*) a water-rich polar phase, most containing of iced H₂O, CO₂,
28 NH₃ in direct contact with the silicate/carbonaceous core, and *ii*) an apolar phase, comprising most
29 of the iced CO, the remaining CO₂ and probably most of the iced CH₃OH.
30
31
32
33
34
35
36
37
38
39
40

41 Investigating iCOMs has been carried out by means of the usual multidisciplinary approach applied
42 in astrochemistry: astronomical spectroscopic observations, astrochemical models, and laboratory
43 experiments. Spectroscopic observations can detect iCOMs in different astronomical sources and
44 provide abundances in the different environments. However, they are not capable to give direct
45 information on how iCOMs are formed, either on the grain-surfaces or in gas-phase. Astrochemical
46 models are useful in rationalizing iCOMs observations. In these models, however, the energetic
47 parameters introduced as input data are associated with some large and critical uncertainties (*i.e.*, in
48 some cases they are derived from gas phase or empirical estimates) and, accordingly, predictions
49 are uncertain too. As a matter of fact, current “grain-surface-formation” models are not capable to
50
51
52
53
54
55
56
57
58
59
60

1
2
3 reproduce the recent observations for methanimine and of methoxymethanol ($\text{CH}_3\text{OCH}_2\text{OH}$), where
4
5 discrepancies of several orders of magnitude were reported.^{6,24} Laboratory experiments are very
6
7 useful in telling us the nature of the products formed but they are not able to reproduce realistically
8
9 the physical conditions of the ISM (*e.g.*, the very low temperature and gas densities or the relatively
10
11 high UV photon or H-atoms fluxes), as well as the chemical features of the ice grains such as the
12
13 exact chemical composition.²⁵

14
15
16 Within this context, computational chemistry is a complementary tool to the other approaches as it
17
18 can alleviate part of the abovementioned problems. Computational simulations can furnish the
19
20 atomistic and electronic structures of the systems under investigation, providing unique information
21
22 such as structural, energetic and spectroscopic features of the ice mantles. They can also provide a
23
24 molecular description of the elementary steps involved in a grain surfaces reaction (*i.e.*,
25
26 adsorption/accretion, diffusion, chemical reactions and desorption) in which a full characterization
27
28 of the energy profiles can be simulated. Interestingly, with these profiles, relevant energetic
29
30 information of the grain surface process (*e.g.*, energy barriers, reaction rates, binding/desorption
31
32 energies) can be obtained, which in turn can be used as accurate input data in the astrochemical
33
34 models. However, this approach also holds some disadvantages: the main one is that results depend
35
36 on both the method chosen to solve the equations describing the systems and the atomistic model
37
38 adopted to represent the grain surface structure. In gas phase calculations this latter disadvantage is
39
40 avoided, and hence different works dealing with the formation of iCOMs through gas phase
41
42 processes are available in the literature, reporting accurate energy profiles and reaction rate
43
44 coefficients.^{26–33}

45
46
47 The very first computational chemistry works dealing with astrochemical reactions on ice surfaces
48
49 date from the beginning of this century. However, the structural ice models were based on the
50
51 presence of a limited number of H_2O molecules or implicit solvation models. It was not since the
52
53 beginning of this decade that ice models started to be structurally more realistic, as periodic slab
54
55
56
57
58
59
60

1
2
3 models, amorphous systems, *etc.* Several works have covered the formation of interstellar
4 molecules on grain surfaces. Some of them addressed the formation of simple compounds on dust
5 grains, (*e.g.*, H₂ and H₂O formation on silicates^{34–37}), but most of them are related to the formation
6 of iCOMs on ice mantles. The aim of the present work is to review all these later studies present in
7 the literature, which to the best of our knowledge is hitherto fully missing. We only focus on
8 computational works based on quantum (QM) or classical (MM) mechanics simulations, as these
9 techniques are the most reliable ones to tackle iCOMs formation.

10
11
12
13
14
15
16
17
18
19
20 The review is organized as follows. Section 2 provides a description of the computational
21 framework, focusing briefly on the quantum chemical methods and techniques, ways to calculate
22 rate constants (including tunneling effects) and ice surface modelling strategies. Section 3 is the
23 core of the review, in which the most relevant computational works dedicated to the “on-surface”
24 iCOM formation are briefly exposed. Here, we are not limited to iCOMs according to the definition
25 given above, but to also other related species, such as H₂CO, CH₃OH, and amino acids and
26 nucleobases. Finally, Section 4 provides the conclusions including some future perspectives in the
27 simulation of iCOMs formation on ice mantles by means of computational chemistry tools.

39 **2 Computational Framework**

40 **2.1 Quantum mechanical methods**

41
42
43
44
45
46
47
48
49
50
51
52
53
54
55
56
57
58
59
60
The chemical processes reviewed here concern a wide variety of reaction-types and mechanisms
(proton/electron transfers, nucleophilic/electrophilic attacks, *etc.*) occurring at structural models
mimicking the surfaces of interstellar ice mantles. Accordingly, accuracy of the results partly relies
on the QM methodologies describing the chemical reactions.

When high accuracy is needed, approaches based on the improvement of the wavefunction such as
the Møller-Plesset 2nd order perturbation theory and coupled cluster single, double and perturbative-
triple electronic extractions method (namely, MP2 and CCSD(T), respectively) are adopted.^{38,39}

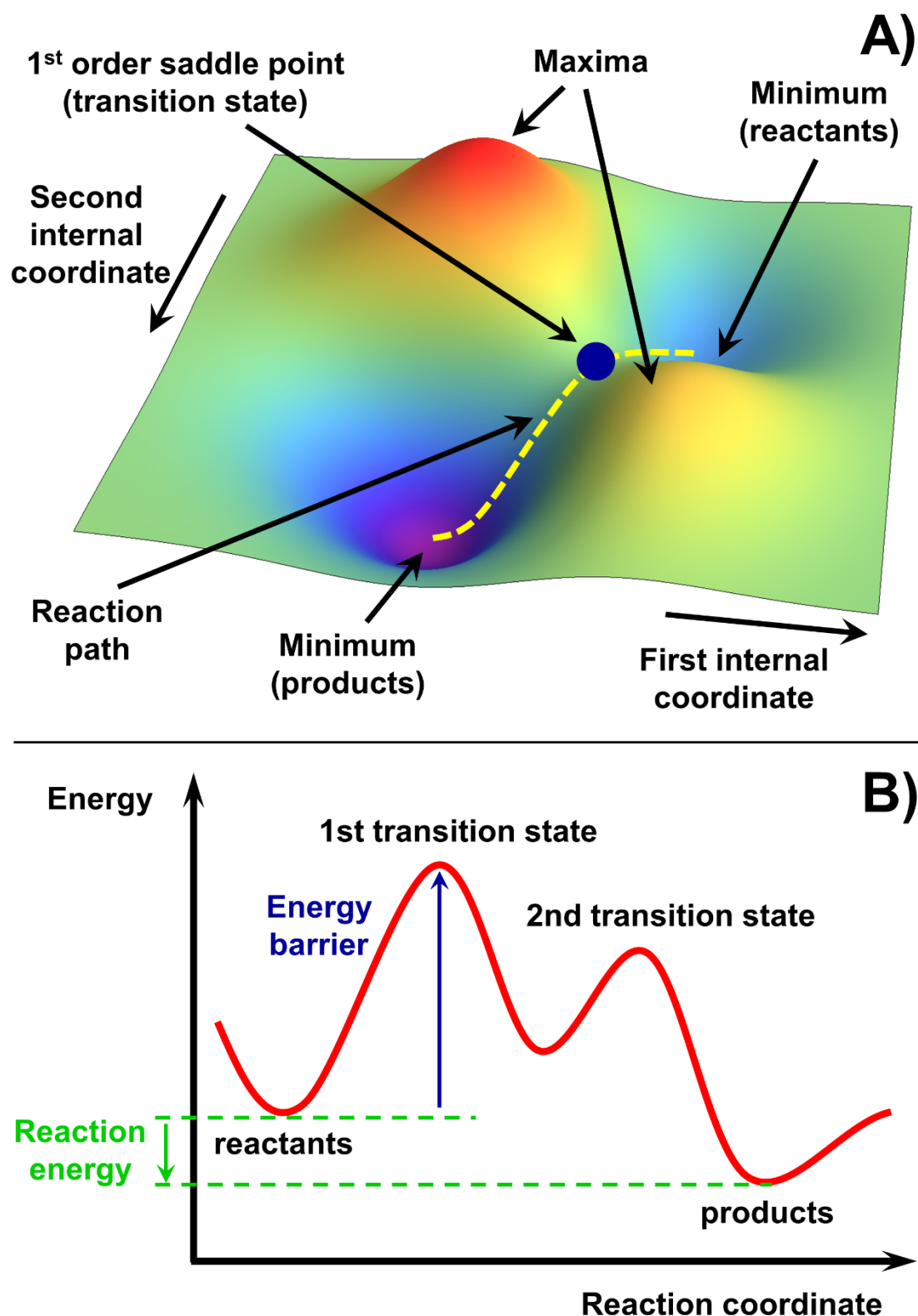
1
2
3 However, these methods are extremely expensive for large systems and accordingly unpractical
4 hitherto when modeling on-surface reactions. Alternatively, since the late 1990s, approaches based
5 on the electron density, the so-called density functional theory (DFT) methods, have become
6 computationally cheaper alternatives to the wavefunction-based ones, in which well-designed DFT
7 methods provide acceptable accuracy.⁴⁰⁻⁴⁴ Among them, the B3LYP, PBE and BHLYP functionals
8 are three the of the most adopted DFT methods in QM calculations.
9
10
11
12
13
14
15

16 17 **2.2 Potential Energy Surfaces**

18
19
20 Potential energy surfaces (PESs) describe the energy of a system (collection of atoms) as a function
21 of its geometry (the position of the atoms). Complete PESs are characterized by calculating the
22 energy of the system as a function of the internal coordinates (bonds, angles and dihedrals).
23 Stationary points (points with a zero gradient in the PES) have physical meaning (Figure 1A):
24 minima correspond to physically stable chemical species (reactants, products and intermediates),
25 while 1st order saddle points correspond to transition states (TSs), the highest energy points on the
26 reaction coordinates (the lowest energy paths connecting reactants with products). All other
27 stationary points (*i.e.*, higher order saddle points and maxima) are physically unsounded. When the
28 PESs are described as a function of the reaction coordinate (the coordinate governing the reaction),
29 the surface is called energy profile (see Figure 1B).
30
31
32
33
34
35
36
37
38
39
40
41
42

43 The nature of the stationary points can be known by diagonalizing the Hessian matrix of second
44 derivatives of the potential energy with respect to the atomic positions. Hessian eigenvalues are
45 related to the frequency vibrational modes of the system: for minima structures, all frequencies are
46 real, while saddle points have one imaginary frequency.
47
48
49
50

51 Since QM calculations account for the electronic structure of the systems, exploration of PESs for
52 reactions (in which chemical bonds break and form) has to be done within this framework (at
53 variance with classical mechanics, which do not account for electrons).⁴⁵⁻⁴⁷
54
55
56
57
58
59
60



53
54
55
56
57
58
59
60

Figure 1. A) Example of a potential energy surface (PES) described as a function of two internal coordinates. The different stationary points are also shown: minima (reactants and products), 1st order saddle point (connecting the minima with the reaction path) and maxima (with no physical meaning). B) Example of an energy profile, in which the PES is described as a function of the reaction coordinate, with the different stationary points. The intrinsic energy barrier (in blue) and the reaction energy (in green) are also shown.

2.3 Static calculations versus dynamic simulations

Exploration of PESs is done systematically by solving the electronic Schrödinger equation for the different stationary points. These are called static calculations because dynamic effects inferred by temperature are not accounted for, *i.e.*, calculations are performed considering 0 K.

Dynamic simulations (also known as molecular dynamics simulations, MDs) allow studying of the evolution in time-space phase of the atomic positions subject to the internal forces of chemical nature and to the kinetic energy due to the temperature of the system. MDs simulations combining electronic structure theory (for electron description) with classical molecular mechanics (for the nuclei motion) are usually referred as *ab initio* molecular dynamics simulations (AIMDs). MDs are adaptable to very different situations: indeed, they can be used to transform a crystalline system into an amorphous one,^{23,48} to sample the adsorbate/surface PES,⁴⁹ to study the diffusion properties of the adsorbates,⁵⁰ or to explore the role of temperature and pressure in surface phenomena.^{51,52}

2.4 Kinetics

Reaction kinetics refers to the velocity of chemical reactions, which are quantified by the rate constant. Quantitatively, reaction rates can be derived from the classical “transition state theory” originally developed by Henry Eyring, Meredith G. Evans and Michael Polanyi in 1935.⁵³ Starting from the assumption of the existence of a “*quasi-equilibrium*” between reactants and TSs, the kinetic rate constant k of a given reaction can be derived as:^{54,55}

$$k = \kappa \frac{k_B T}{h} e^{-\frac{\Delta G^\ddagger}{RT}} (c^0)^{1-m} \quad \text{Eq. 1}$$

where κ is the transmission coefficient (for reactions without tunneling assumed to be 1), T the absolute temperature, k_B the Boltzmann constant, h the Planck constant, R the ideal gas constant, c^0 the standard concentration, m the molecularity ($m = 1$ or 2 for uni or bimolecular reactions) and ΔG^\ddagger the Gibbs free energy barrier, *i.e.* the free energy difference between the TS and the reactants.

For a unimolecular reaction ($m = 1$), k can be easily related to the half-life time $t_{1/2}$;⁴⁷ *i.e.*, the time needed to consume the half of the initial amount of reactants:

$$t_{1/2} = \frac{\ln 2}{k} = \frac{h}{\kappa k_B T} e^{\frac{\Delta G^\ddagger}{RT}} \ln 2 \quad \text{Eq. 2}$$

At the very low temperatures of the ISM long half-life times are derived, even for very low energy barriers. The dependence of $\log(t_{1/2}/1 \text{ Myr})$ on ΔG^\ddagger for different typical temperatures of ISM is reported in Figure 2. Data shown in the inset clearly indicate that, in the 10-25 K temperature range of MCs,⁵⁶ only reactions with very low kinetic barriers can occur.

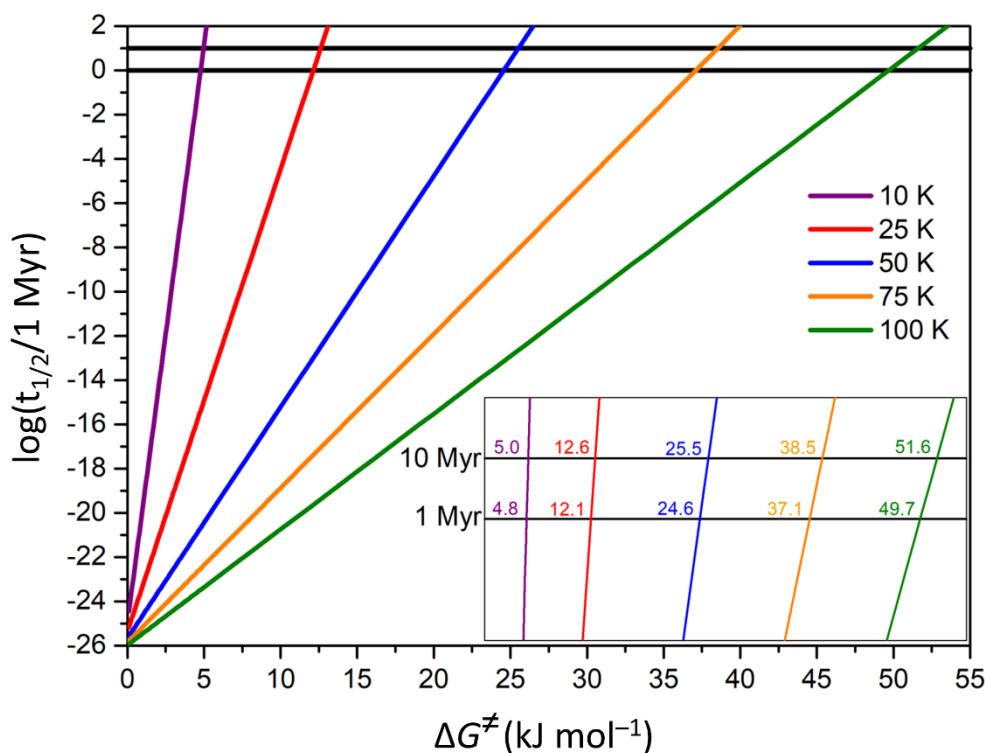


Figure 2. Dependence of the half-life time ($t_{1/2}$) with the free energy barrier (ΔG^\ddagger) for a unimolecular reaction at different temperatures (in K). $t_{1/2}$ are normalized to 1 Myr. The two straight horizontal lines represent 1 and 10 Myr, respectively, taken as reference lifetimes of a typical molecular cloud.⁵⁷ Inset: zoomed view in the $-2 \leq \log(t_{1/2}/1 \text{ Myr}) \leq 2$ range. Numbers at the crossing points are the ΔG^\ddagger values at which $t_{1/2}$ equals 1 and 10 Myr, respectively.

It is worth mentioning that at the particularly low interstellar temperatures, and for not too high and wide barriers, quantum tunneling may play a prominent role favoring reaction rates. There are several ways to account for such tunneling effects,⁵⁸ such as the semi-classical approaches, in which

1
2
3 the the transmission coefficient κ is calculated through specific formulae (*e.g.*, the Eckart formula,⁵⁹
4 usually used in astrochemical modeling). More evolved is the instanton theory,^{60,61} which is a
5 derivation of the harmonic quantum transition state theory,⁶² where the tunneling path is fully
6 optimized in the Feynman-path-based instanton theory.⁶³ Few examples on the use of the instanton
7 theory in astrochemical reactions can be found elsewhere.^{64,65}

8
9
10 The free energy of a given species can be easily obtained once computed the partition functions
11 (translation, rotational, electronic and vibrational) and applying statistical thermodynamics
12 relations.⁶⁶ The free energy of a species at temperature T is given by:

$$G(T) = E + \zeta + \epsilon(T) + PV(T) - TS(T) \quad \text{Eq. 3}$$

13
14
15 where E is the potential energy (electronic plus nuclear) of the species from the electronic structure
16 calculation, ζ is the zero-point energy (ZPE), $\epsilon(T)$ is the thermal contribution to the internal energy
17 (both terms obtained with frequency calculations), and P , V and S represent the volume, pressure
18 and entropy, respectively. At the low temperatures of the ISM, the last three terms of Eq. 3 are
19 small and usually neglected. Thus, energy profiles are usually presented in terms of E or $E + \zeta$,
20 with the latter being referred to as “internal energy at 0 K” or, equivalently, “enthalpy at 0 K”.

2.5 Surface Modeling

21
22
23 Accuracy of the theoretical results, in addition to the QM methods (see above), also relies on the
24 specific models adopted to simulate the ice surfaces. Two strategies can be adopted to model the
25 external surfaces of icy grains: the periodic approach and the cluster approach. The periodic
26 approach consists of applying the periodic boundary conditions (PBC) into a unit cell containing the
27 surface adsorptive/catalytic sites, resulting in an infinite 2D slab model, *i.e.*, periodicity is only
28 applied in the two directions defining the unit cell (Figure 3A).^{23,67–77} In contrast, the cluster
29 approach consists of cutting out from the periodic model a finite set of atoms containing the surface
30 sites, so that the surface is essentially modelled by a molecular system (Figure 3B).

1
2
3 Powerful computer codes have been developed over the years to solve the PBC problem for infinite
4 systems. However, due to their infinite nature, application of highly accurate wave function-based
5 methods is overwhelming (although recent developments indicate applicability for MP2⁷⁸) and they
6 can in practice only be studied using DFT methods. Moreover, localization of transition state
7 structures is less developed compared to molecular codes, thus PES characterization being limited
8 to “simple” reactions. On the contrary, a large variety of quantum molecular programs can properly
9 handle cluster models, characterizing PESs of complex chemical reactions, using even CCSD(T),
10 depending on the cluster size.
11
12
13
14
15
16
17
18
19
20

21
22 The cluster approach can be limited by: *i*) the need to “heal” dangling bonds resulting from cutting
23 covalent/ionic bonds from the extended system, and *ii*) the size of the cluster, which should be large
24 enough to include the catalytic sites. For this latter, cluster sizes can prohibitively be large, reducing
25 the abovementioned advantages when adopting molecular computer codes. A possible solution is to
26 use embedding techniques like the ONIOM method,^{79–81} in which the region of interest (*e.g.*, the
27 region close to the catalytic sites) is treated at high level of theory (MP2, CCSD(T)), whereas the
28 surrounding region is treated at a lower level (DFT, semi-empirical or even molecular mechanics,
29 Figure 3C).
30
31
32
33
34
35
36
37
38
39

40
41 Interstellar ices are usually reported to be highly amorphous and, partly, porous,^{20,82–84} although the
42 degree of porosity has recently been questioned.²² Amorphous surfaces (Figure 3D) can be
43 generated by amorphizing (*e.g.*, running MDs at high temperature) the slab model, or by cutting out
44 a previously amorphized bulk system. The presence of pores can influence the reactivity on
45 interstellar ices since: *i*) adsorbates can be entrapped and retained inside the pore (hence favoring
46 reaction with other entrapped species), and *ii*) water molecules may exert a “solvent-like” effect,
47 thus stabilizing intermediates or transition states. A consistent way to simulate pores is through
48 clathrate models, as clathrate-like IR features have been identified in interstellar ices,⁸⁵ and the
49 presence of different H₂O-clathrate encapsulated species in Earth’s⁸⁶ and Titan’s atmospheres.⁸⁷
50
51
52
53
54
55
56
57
58
59
60

1
2
3 Interestingly, even for amorphized systems, ice water molecules tend to form clathrate-like cages.²³
4
5 However, to the best of our knowledge, no theoretical works addressing iCOMs formation using
6
7 clathrate atomistic models are available, while those focusing on the clathrate-molecule interactions
8
9 are scarce.^{23,88} A way to account for the “pore stabilizing effects” is by using the “polarizable
10
11 continuum model” (PCM).^{89,90} PCM is a computationally cheap technique in which solvation
12
13 effects are described with a continuous dielectric constant ϵ (the value of liquid water, 78.5, is
14
15 usually used to simulate solid water⁹¹⁻⁹³). Reactive compounds are immersed within the continuum
16
17 dielectric medium (Figure 3E). However, as solvent molecules are not explicitly considered,
18
19 specific ice-molecule interactions are omitted. Although this can partly be solved by introducing a
20
21 “first hydration sphere” of explicit water molecules within the PCM cavity,⁹⁴ using a reduced
22
23 number of water molecules without geometrical constraints can convert the initial pore into a
24
25 surface due to aggregation phenomena between water molecules.²³
26
27
28
29
30
31
32
33
34
35
36
37
38
39
40
41
42
43
44
45
46
47
48
49
50
51
52
53
54
55
56
57
58
59
60

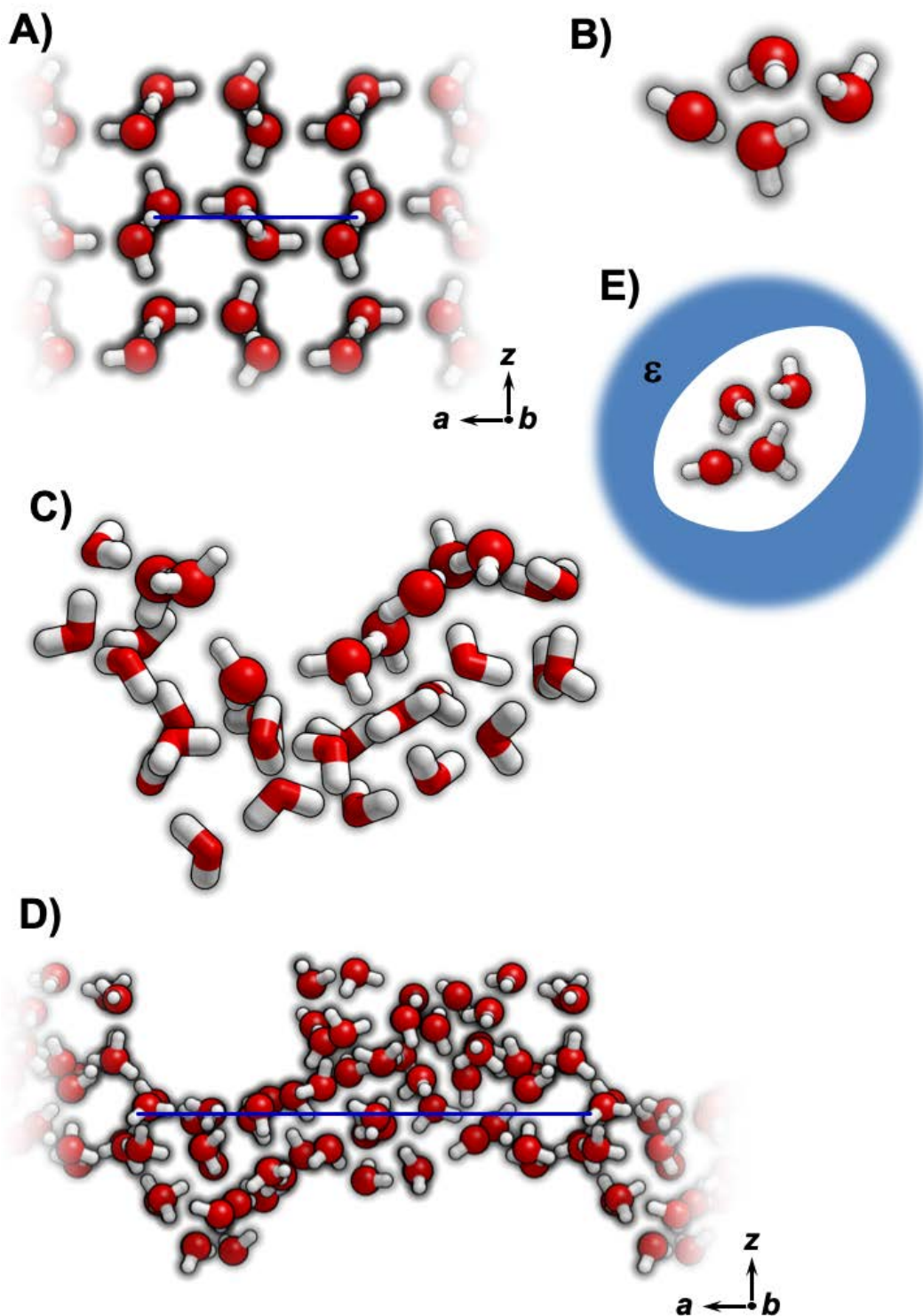


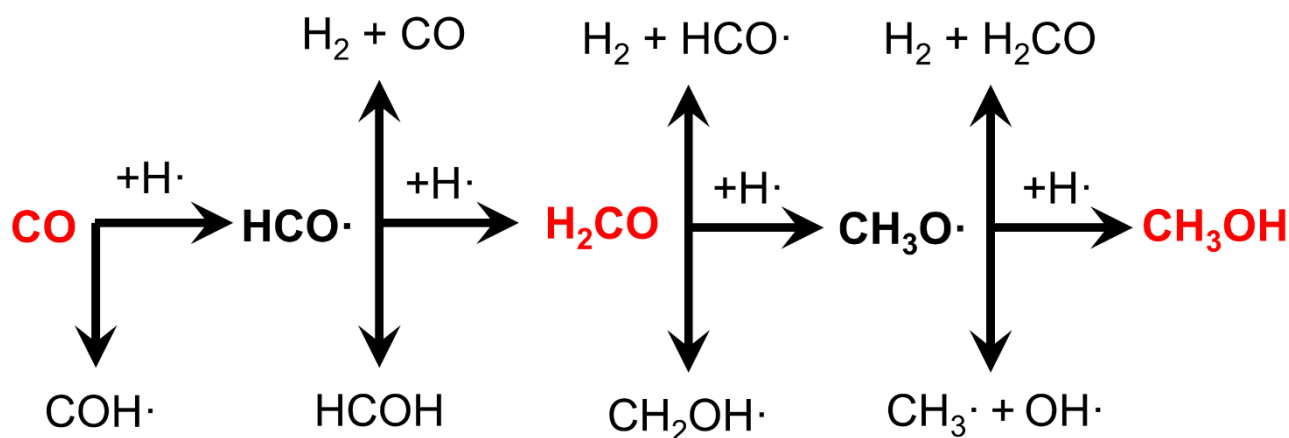
Figure 3. Different strategies to model water ice surfaces: A) Side view of a crystalline 2D-periodic slab model; B) Minimal cluster of 4 waters; C) ONIOM approach for a 33 H_2O cluster: molecules as balls represent the “high level”, those as stick the “low level”. D) Side view of an amorphous 2D-periodic slab model. E) PCM approach for a 4 H_2O cluster (blue background represents the continuum dielectric ϵ). For A) and C), a and b are the periodic vectors and z is the non-periodic direction (the a vector is represented in blue). H-bonds among water molecules are not represented. Colour legend: oxygen in red, hydrogen in white.

3 Computational chemistry works for iCOMs formation

3.1 Formaldehyde (H₂CO) and methanol (CH₃OH) formation

Formaldehyde (H₂CO) and methanol (CH₃OH) are among the few molecules that have been widely detected as components of the icy mantles.²⁰ From the point of view of iCOMs formation, these two compounds are very important because they are the precursors of more complex species. For instance, their dissociation leads to the formation of HCO·, CH₃O· and CH₂OH· radicals, which can trigger reactions forming iCOMs.

Surface formation of H₂CO and CH₃OH, firstly postulated⁹⁵ and then confirmed experimentally,^{14,96} takes place through successive hydrogenation of CO, which was previously accreted onto dust grains (see Scheme 1, horizontal path). However, these reactions present competitive processes (represented by the vertical paths), which can make less efficiency H₂CO and CH₃OH formation.



Scheme 1. Formation of formaldehyde (H₂CO) and methanol (CH₃OH) from successive H·-additions to carbon monoxide (CO, horizontal path). Vertical paths refer to competitive channels. Adapted from Ref. 97.

David E. Woon computed the PESs of the first (CO + H· → HCO·) and third (H₂CO + H· → CH₃O·) hydrogenation in the presence of (H₂O)_n clusters ($n = 0-4$ and 12). For $0 \leq n \leq 4$, post-Hartree-Fock calculations were performed, which were complemented with the PCM solvation approach.⁹⁷ Calculated energy barriers varied from 17-70 kJ mol⁻¹, depending on the QM method, n , and PCM application or not. Author concluded that water molecules did not possess a specific

1
2
3 catalytic role in the reactions, probably playing an indirect role (*e.g.*, third body), and opening the
4 possibility that tunneling effects could be important for their occurrence.
5
6

7
8 More recently, Rimola *et al.*⁹⁸ simulated the same reactions in gas phase and in the presence of ice
9 surfaces made up by 3, 18 and 32 H₂O molecules at the BHLPY DFT level. Authors indicated that
10 both reactions presented exceedingly high energy barriers ($\approx 9\text{-}14\text{ kJ mol}^{-1}$) to occur at 10-20 K.
11
12 Accordingly, tunneling effects were advocated for the occurrence of the reactions. Despite this,
13 authors underlined a catalytic role of water ice since on the ice the energy barriers were slightly
14 lower than in gas phase. Such catalytic effects were associated with bond polarizing effects caused
15 by the interaction of CO and H₂CO with H₂O surface ice molecules, *i.e.*, the C–O bonds became
16 weakened upon interaction, making the C atom more prone to be hydrogenated. Similar results were
17 also found by Goumans *et al.*^{99,100} when the reactions occurred on hydroxylated silica surfaces.
18
19 Here, the C–O bond polarization was induced by the surface Si–OH groups.
20
21

22
23 Finally, Woon identified an alternative “on-ice” synthetic route for CH₃OH.¹⁰¹ In this work, it was
24 found that interaction of CH₃⁺ with H₂O ice led first to the formation of CH₃OH₂⁺ (*i.e.*, protonated
25 methanol) and then to the release of the extra proton to the ice to finally form CH₃OH, *i.e.*, CH₃⁺ +
26 (H₂O)_n → CH₃OH₂⁺ + (H₂O)_(n-1) → CH₃OH + H₃O⁺ + (H₂O)_(n-2) (see Figure 4). All these processes
27
28 were found to be barrierless, *i.e.*, they occurred spontaneously during the geometry optimization.
29
30 Despite the novelty of the path, authors highlighted its dependence on the CH₃⁺ interstellar
31
32 abundance, a controversial aspect since direct observation of CH₃⁺ is difficult due to transition
33
34 symmetry rules so tentative detections are complemented with its CH₂D⁺ isotopolog.^{102,103}
35
36
37
38
39
40
41
42
43
44
45
46
47
48
49
50
51
52
53
54
55
56
57
58
59
60

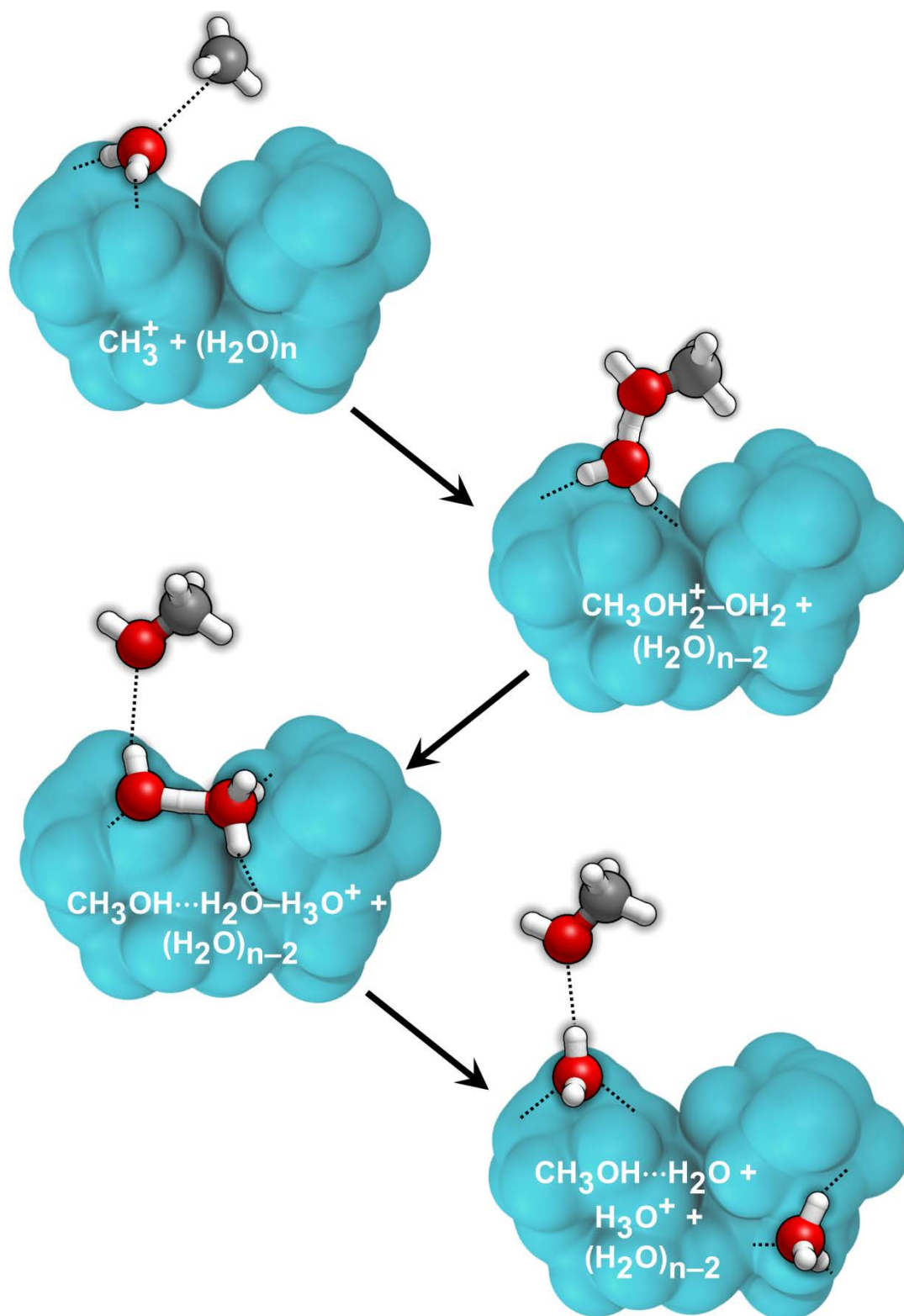


Figure 4. Schematic representation of CH₃OH formation from CH₃⁺. Some H₂O molecules are explicitly shown, while the rest are rendered in light blue. Adapted from Ref. 101. Colour legend: oxygen in red, carbon in grey, hydrogen in white.

3.2 Formamide (NH₂CHO) formation

Formamide (NH₂CHO) is one of the molecules that attracted great attention in the last years. It was first detected in 1971 in the massive star forming regions Sgr B2 and in Orion KL1,⁷ and since then dedicated observational campaigns have revealed its presence in a variety of star-forming regions, shock sites and protostellar objects,^{104–109} as well as comets,^{110,111} suggesting a relatively widespread abundance. The astrochemical relevance of formamide arises from manifold aspects: *i*) it is the simplest iCOM containing the four most essential elements for biological systems (*i.e.*, H, C, N and O), *ii*) it is the simplest organic compound containing the amide bond –C(=O)–NH–, the same bond joining amino acids into peptides, and *iii*) there is experimental evidence that it is an effective reactant for the synthesis, in the presence of naturally-occurring minerals and oxides, of precursor biomolecules constituting metabolic and genetic material (see Section 3.8).

Formamide, as other iCOM, is not exempted from the debate whether its formation occurs in the gas phase or on the surfaces of the icy grain mantles. Several theoretical works addressed its gas phase formation through diverse ion-molecule reactions^{32,112} and the bimolecular reaction of H₂CO + NH₂· → NH₂CHO + H·.^{27,28} On the ice mantles, its formation has also been addressed by several authors.

Song and Kästner studied the HNCO hydrogenation (*i.e.*, H· + HNCO → NH₂CO·) on an amorphous water ice cluster model at a hybrid QM/MM theory level.⁶⁴ The second hydrogenation was considered to be barrierless, involving a radical-radical reaction. This synthetic route was studied in view of the linear correlation between NH₂CHO and HNCO abundance in different sources.¹⁰⁶ On the ice surfaces, calculated energy barrier adopting an Eley-Rideal mechanism was found to be 4 kJ mol⁻¹ lower than in gas phase (31.8 and 36.2 kJ mol⁻¹ respectively) due to bond polarizing effects exerted by the ice (similar to hydrogenation of CO and H₂CO, see above⁹⁸). However, tunneling rate constants obtained with the instanton theory were found to be low at the low ISM temperatures, due to a tunneling inefficiency caused by the broad energy barrier width.

1
2
3 These results were in agreement with the inefficient hydrogenation of HNCO ices found
4 experimentally.¹¹³
5
6

7
8 Another reaction channel investigated theoretically is the $\text{NH}_2\cdot + \text{HCO}\cdot \rightarrow \text{NH}_2\text{CHO}$ radical-radical
9 coupling on a $(\text{H}_2\text{O})_{33}$ cluster model at BHLYP DFT theory level.¹¹⁴ This “simple” radical coupling
10 followed the usual scheme proposed for iCOMs formation in several astrochemical models,^{11,115}
11 and tested experimentally.¹¹⁶ Results indicated that the actual biradical system (*i.e.*, the two radicals
12 adsorbed on the ice surface with opposite spin states) was stable, precisely because of the
13 interaction with the surface, and that the coupling had an energy barrier of 3 kJ mol^{-1} . However, it
14 was also found that, when the two radicals are properly oriented, a direct H \cdot -transfer from HCO \cdot to
15 NH $_2\cdot$ leading to CO + NH $_3$ occurred in a barrierless way. H \cdot -transfers of this kind were also
16 observed in acetaldehyde (CH $_3$ CHO) formation HCO \cdot + CH $_3\cdot$ by Enrique-Romero *et al.*:¹¹⁷ in this
17 case, CH $_3$ CHO formation competed with CO + CH $_4$ formation, pointing out that reactivity between
18 radicals not always leads to iCOMs formation.
19
20
21
22
23
24
25
26
27
28
29
30
31
32

33
34 In the same work of Rimola *et al.*,¹¹⁴ two additional synthetic paths were investigated: reaction of
35 one H $_2$ O molecule of the ice with either HCN or CN \cdot . The first path (*i.e.*, $\text{H}_2\text{O} + \text{HCN} \rightarrow$
36 NH_2CHO) was found to have very large energy barriers (167 kJ mol^{-1}), and therefore unfeasible in
37 ISM. The second one (*i.e.*, $\text{H}_2\text{O} + \text{CN}\cdot \rightarrow \text{NH}_2\text{CO}$), however, was found to be energetically
38 favorable due to two aspects: *i*) the high reactivity of the CN \cdot radical, and *ii*) water ice acts as a
39 catalyst by lowering the energy barrier. Indeed, here we present one of the most important aspects
40 of water in the reactions of iCOMs formation, *i.e.* its capability to act as a hydrogen-transfer
41 assistant, with hydrogen having a proton (H $^+$) character. Upon this role, water molecules belonging
42 to the ice exchange H $^+$, *i.e.*, they receive one H $^+$ releasing at the same time another one, helping the
43 transfer process. This role of H $^+$ -transfer assistant can be shared by different water molecules, thus
44 establishing a H $^+$ relay mechanism. Such a behavior allows both the occurrence of H $^+$ -transfers
45 through a chain of well-connected water molecules and the reduction of the geometrical strains in
46
47
48
49
50
51
52
53
54
55
56
57
58
59
60

1
2
3 TS structures with respect to the gas phase, hence stabilizing them and lowering the energy barriers
4
5 of the associated H^+ -transfer process. As an example, the TS structures for the $HNCOH\cdot \rightarrow$
6
7 $NH_2CO\cdot$ isomerization for water ice acting as H^+ -transfer assistant and in gas phase are reported in
8
9 Figure 5A and B, respectively: the strongly geometrical-strained four-member ring in the gas phase
10
11 becomes a low strained eight-member ring when three water molecules are present. The final step
12
13 leading to the formation of the actual NH_2CHO from $NH_2CO\cdot$ was proved to occur via either $H\cdot$ -
14
15 addition (barrierless) or *via* $H\cdot$ -abstraction of a H_2O ice molecule, in which kinetic results indicated
16
17 a fast-overall process ($k \sim 10^{-9} s^{-1}$).
18
19
20

21
22 Finally, Bredehöft *et al.*¹¹⁸ studied – by combining experiments and theory – the synthesis of
23
24 NH_2CHO under electron exposure of $NH_3:CO$ ice mixtures. Experiments detected NH_2CHO
25
26 formation and calculations provided a molecular interpretation of these findings (only considering
27
28 the reactive species, namely, without considering the rest of the ice components). The mechanistic
29
30 key point was the formation of a transient radical anion NH_3^- , which triggered the following multi-
31
32 step reaction: *i*) formation of $NH_2\cdot$ and H^- (barrierless), *ii*) reaction of $NH_2\cdot + CO \rightarrow NH_2CO\cdot$
33
34 (barrierless), and *iii*) reaction of $NH_2CO\cdot + NH_3 \rightarrow NH_2CHO + NH_2\cdot$, in which the excess energy
35
36 provided by the electron attachment was advocated to help overcoming the high energy barrier (≈ 65
37
38 $kJ mol^{-1}$).
39
40
41
42
43
44
45
46
47
48
49
50
51
52
53
54
55
56
57
58
59
60

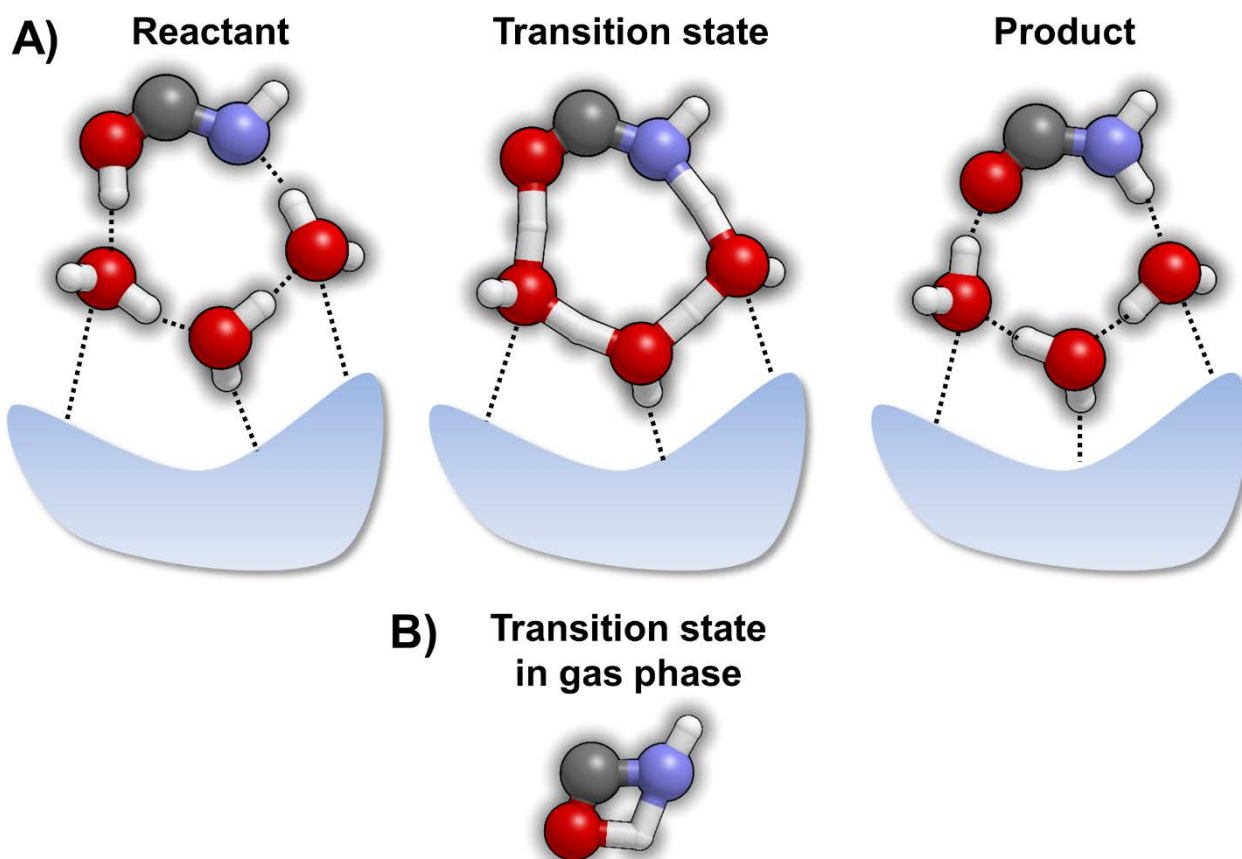


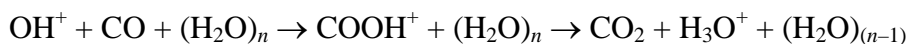
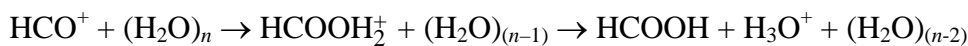
Figure 5. A) Schematic representation of the role of water ice acting as H^+ -transfer assistant. In this case, three water molecules are helping the transfer adopting a relay mechanism for the $\text{HNCOH}\cdot \rightarrow \text{NH}_2\text{CO}\cdot$ isomerization occurring in the $\text{CN}\cdot + \text{H}_2\text{O} \rightarrow \text{NH}_2\text{CO}\cdot$ reaction.¹¹⁴ B) Transition state structure for the $\text{HNCOH}\cdot \rightarrow \text{NH}_2\text{CO}\cdot$ isomerization in gas phase. Colour legend: oxygen in red, carbon in grey, hydrogen in white.

3.3 Acidic species

3.3.1 Formation HCOOH (and related species) and its reactivity

Woon studied the $\text{OH}\cdot$ -addition to CO forming the *trans*- $\text{COOH}\cdot$ radical, whose dehydrogenation led to the formation of CO_2 , *i.e.*, $\text{CO} + \text{OH}\cdot \rightarrow \text{trans-COOH}\cdot \rightarrow \text{CO}_2 + \text{H}\cdot$.¹¹⁹ Formation of *trans*- COOH was found to be more favorable than CO_2 in PCM (energy barriers ≈ 13 and ≈ 40 kJ mol^{-1} , respectively). In the same work, it was shown that the radical-radical coupling between *trans*- $\text{COOH}\cdot$ and $\text{CH}_2\text{NH}_2\cdot$ yielded glycine formation in a barrierless fashion.

Woon also investigated the formation of HCOOH and CO_2 from the precursors of HCO^+ and “ $\text{OH}^+ + \text{CO}$ ”, respectively, in a similar way to CH_3OH formation from CH_3^+ (see Section 3.1),¹⁰¹ *i.e.*:



These processes were identified to occur spontaneously during the geometry optimizations when HCO^+ and “ $\text{OH}^+ + \text{CO}$ ” interacted with water molecules of the ice models.

Rimola *et al.*¹²⁰ studied the $\text{CO} + \text{OH}\cdot \rightarrow \text{COOH}\cdot$ reaction on a cage-like $(\text{H}_2\text{O})_8$ -derivative cluster as water ice surface model. The $\text{OH}\cdot$ reactant was initially formed by processing of the $(\text{H}_2\text{O})_8$ cluster (see Figure 6), *i.e.*, photolytic removal of one H atom, leading to the formation of a radical neutral cluster (RN path in Figure 6), and one electron removal, leading to the formation of a radical cation cluster (RC path in Figure 6) showing both $\text{OH}\cdot$ and H_3O^+ . Reaction of these $\text{OH}\cdot$ species with CO to form $\text{COOH}\cdot$ (Figure 6, last steps) were computed at BHLYP level providing relatively low energy barriers (14 and 12 kJ mol^{-1} , respectively).

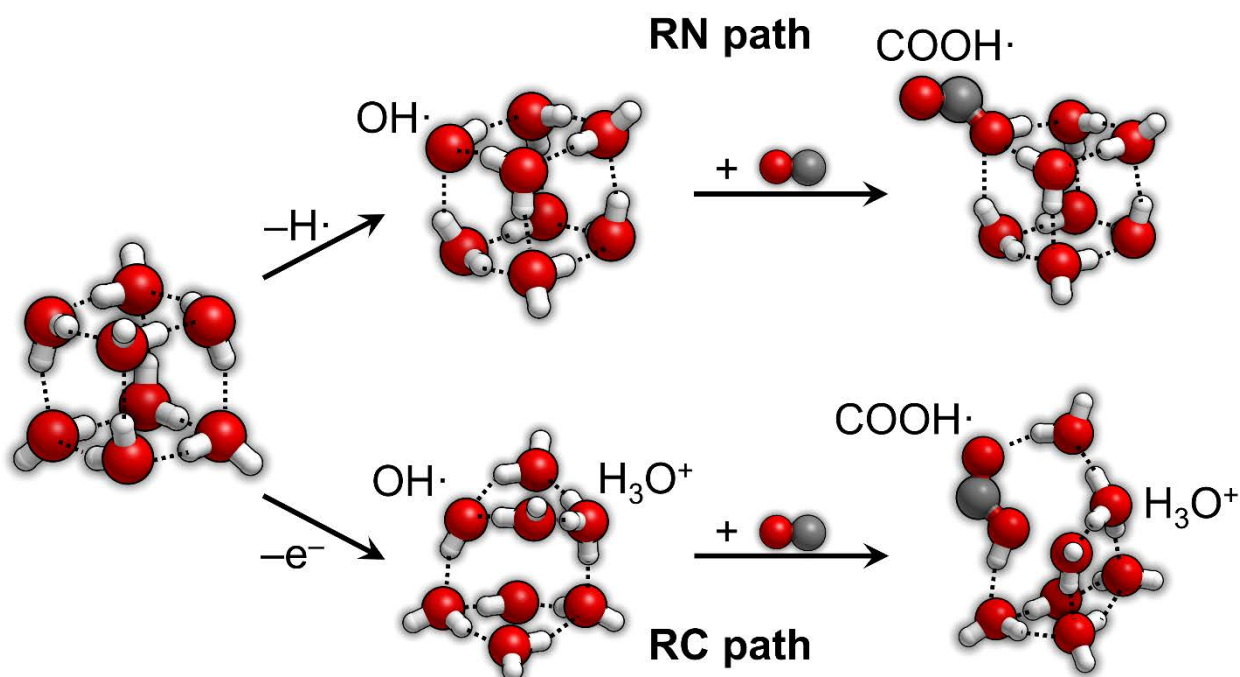


Figure 6. Schematic representation of the formation of $\text{COOH}\cdot$ from $\text{CO} + \text{OH}\cdot$ reaction for the RN and RC paths (see text). $\text{OH}\cdot$ comes from the processing of a $(\text{H}_2\text{O})_8$ cluster. Adapted from Ref. 120. Colour legend: oxygen in red, hydrogen in white, carbon in grey.

In relation to the HCOOH reactivity, Woon studied the amination of HCOOH in the presence of 0, 1 and 2 H_2O molecules at MP2 level and adopting PCM.⁹¹ Reaction of HCOOH with NH_3 led first

1
2
3 to $\text{NH}_2\text{CH}_2(\text{OH})_2$ (*i.e.*, hydrated formamide), which eventually dehydrated to give NH_2CHO . In the
4
5 same work, direct formation of glycine through reaction of HCOOH with $\text{CH}_2=\text{NH}$ (methanimine)
6
7 was also investigated. Examined reactions showed large energy barriers, the lowest one being 49 kJ
8
9 mol^{-1} for the formation of $\text{NH}_2\text{CH}_2(\text{OH})_2$ in presence of two explicit water molecules, while direct
10
11 glycine formation showed a very high energy barrier (406 kJ mol^{-1}). Generally, PCM solvation
12
13 effects lowered the energy barriers by about $5\text{-}45 \text{ kJ mol}^{-1}$, while the largest energy decreases were
14
15 observed when water molecules acted as H^+ -transfer assistants (about 80 kJ mol^{-1} , at the most).
16
17

18
19 Park and Woon focused on the protonation of NH_3 from HCOOH in the presence of explicit waters
20
21 (2-7, 9, 14, 15 molecules) at B3LYP.¹²¹ The purpose of this work was, rather investigating the
22
23 reactivity, to reproduce the IR features of the $\text{HCOO}^-/\text{NH}_4^+$ solid ion pair. In the presence of at least
24
25 3 H_2O molecules, H^+ -transfer was found to be barrierless. Using “large” (*i.e.*, 7, 9, 14 and 15) H_2O
26
27 clusters, simulated vibrational features reproduced fairly well the observed ones.¹²¹
28
29

30
31 Finally, Kayi *et al.* found that CO_2 and methylamine (CH_3NH_2) on $(\text{H}_2\text{O})_n$ ($0 \leq n \leq 20$ clusters)
32
33 formed the $\text{CH}_3\text{NH}_2^+/\text{CO}_2^-$ ion pair as a result of a charge transfer from CH_3NH_2 to CO_2 .¹²² It was
34
35 identified that water favored the charge transfer and the ion pair stabilization.
36
37

39 3.3.2 Reactivity of other acidic species: HOCN/HNCO , HCN/HNC and CH_3COOH

40
41
42 Park and Woon dealt with the deprotonation of cyanic (HOCN) and isocyanic (HNCO) acids
43
44 reacting with NH_3 with the purpose to: *i*) simulate the reactive processes,¹²³ and *ii*) check if the
45
46 “ XCN ” interstellar band could correspond to one of the resulting species.¹²⁴ Different approaches
47
48 were employed to simulate the water environments, *i.e.*, PCM, small water clusters calculated at full
49
50 QM methods, and large water clusters calculated with the ONIOM strategy. The main conclusions
51
52 of these works were that water-assisted deprotonation of both cyanic and isocyanic acid was
53
54 barrierless in water environments, forming the $\text{OCN}^-/\text{NH}_4^+$ ion pair, and that this pair reproduced
55
56 reasonably well the observationally IR features of the “ XCN ” band, suggesting OCN^- as a good
57
58
59
60

1
2
3 candidate carrier. It was also shown that both HOCN and HNCO spontaneously deprotonated even
4
5 in absence of NH_3 , thus forming an $\text{OCN}^-/\text{H}_3\text{O}^+$ ion pair.
6
7

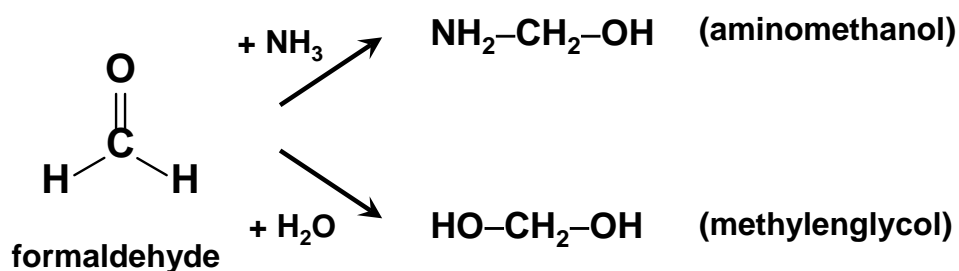
8 Both hydrogen cyanide (HCN) and isocyanide (HNC) have been observed in the ISM,^{125,126} which
9
10 can be interconverted by the $\text{HCN} \leftrightarrow \text{HNC}$ isomerization. There are essentially two works dealing
11
12 with this isomerization in water environments. Garderbien and Sevin investigated the reaction in the
13
14 presence of $(\text{H}_2\text{O})_n$ ($n = 1-4$) clusters at CCSD(T) level.¹²⁷ $\text{HCN} \rightarrow \text{HNC}$ conversion was found to
15
16 be thermodynamically disfavoured, the $\text{HCN}/(\text{H}_2\text{O})_n$ complexes being about 50 kJ mol^{-1} more
17
18 stable than the $\text{HNC}/(\text{H}_2\text{O})_n$ ones. Additionally, calculated energy barriers were 178, 124, 96 and 80
19
20 kJ mol^{-1} for 1, 2, 3 and 4 H_2O acting as H^+ -transfer catalysts, respectively. In the second work,
21
22 Koch *et al.* investigated the $\text{HNC} \rightarrow \text{HCN}$ transformation at B3LYP combining the presence of 3 +
23
24 4 water molecules (representing the first + second hydration spheres) with PCM.⁹⁴ Authors
25
26 identified a progressive energy barrier decrease when adding successively solvation effects (*i.e.*,
27
28 first hydration sphere, the second one, and PCM), reaching the lowest free energy barrier of 11 kJ
29
30 mol^{-1} at 50 K (all solvation effects accounted for), leading to a half-life time of 714 s. The
31
32 discrepancies between this two works can be definitely assigned to different adopted models and
33
34 methods.
35
36
37
38
39

40
41 Finally, Woon investigated the protonation of NH_3 with acetic acid (CH_3COOH) and HCN/HNC in
42
43 presence of $(\text{H}_2\text{O})_n$ ($n = 2-6$) clusters at B3LYP and MP2 levels.¹²⁸ For all the considered processes,
44
45 protonation of NH_3 *via* water-assisted mechanisms became barrierless when at least the $(\text{H}_2\text{O})_3$
46
47 cluster was considered.
48
49

50 51 **3.4 Aminoalcohols formation**

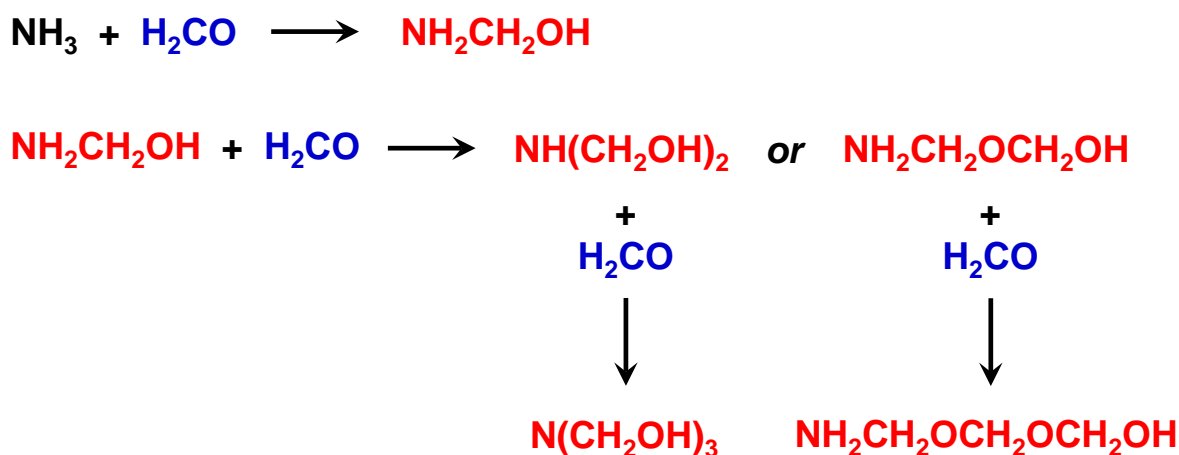
52
53 Aminoalcohols are organic compounds containing both the alcohol ($-\text{OH}$) group and the amino
54
55 groups, this latter being primary ($-\text{NH}_2$), secondary ($-\text{NH}$) or tertiary ($-\text{N}$).
56
57
58
59
60

Addition of NH_3 to H_2CO yields the formation of aminomethanol ($\text{NH}_2\text{CH}_2\text{OH}$), the simplest aminoalcohol. However, this process in water environments has methyleneglycol (HOCH_2OH) formation as a competitive channel (see Scheme 2). In both cases, reactions proceed through a nucleophilic attack of $\text{NH}_3/\text{H}_2\text{O}$ to the C atom of H_2CO followed by a proton transfer to the O atom of the aldehyde $\text{C}=\text{O}$ group.



Scheme 2. Formation of aminomethanol and methyleneglycol by addition of NH_3 and H_2O , respectively, to formaldehyde.

In a seminal work, Woon investigated these two reactions, alongside polymerization of H_2CO as a possible by-side process, in the presence of explicit water molecules and the PCM solvation model at MP2 level.¹²⁹ $\text{NH}_2\text{CH}_2\text{OH}$ formation was found to be the process with the lowest energy barrier (2.5 kJ mol^{-1}) in detriment of HOCH_2OH and H_2CO -polymer formations (energy barriers of 13 and 300 kJ mol^{-1} , respectively). The same author published another article dealing with the formation of more complex aminoalcohols by successive reaction of H_2CO with aminoalcohols formed in previous steps, as well as polyoxymethylenamine $\text{H}-(\text{-OCH}_2\text{-})_n\text{-NH}_2$ (POM- NH_2), an aminoalcohol polymeric form (Scheme 3).¹³⁰ Results indicated that, regardless of the catalytic effects exerted by the PCM environment and the presence of explicit water molecules, these reactions were not likely to occur in interstellar conditions because of their relatively high energy barrier, with values ranging from 20 to 140 kJ mol^{-1} .



Scheme 3. Successive reactivity of formaldehyde (in blue) with aminoalcohols (in red).

Courmier *et al.*¹³¹ refined Woon's calculations¹²⁹ for the $\text{NH}_2\text{CH}_2\text{OH}$ formation in the presence of $(\text{H}_2\text{O})_n$ ($n = 0-3$) explicit molecules at CCSD(T) level. Figure 7A shows the initial structure of the reaction, in which NH_3 and H_2CO interacted with a the $(\text{H}_2\text{O})_3$ cluster adopting a pentamer-like configuration in the way to maximize the H-bond interactions. All calculated energy barriers were found to be systematically higher by about 15 kJ mol^{-1} than those obtained by Woon.¹²⁹ Nevertheless, the role of water acting as H^+ -transfer assistant was clearly shown, with energy barriers of 150, 75 and 50 kJ mol^{-1} for $n = 0, 1$ and 2, respectively. For $n = 3$ the energy barrier was found to slightly increase compared with $n = 2$.

In a more recent work, Rimola *et al.*¹³² simulated the same reaction of $\text{NH}_2\text{CH}_2\text{OH}$ formation on a water ice surface modelled by 18 H_2O molecules at B3LYP level. Figure 7B shows the initial structure of the reaction. An energy barrier of 40 kJ mol^{-1} was computed, disagreeing by some amount with that of 60 kJ mol^{-1} found by Courmier *et al.*¹³¹. The reasons of such a difference arise from both the different theory levels and the different water ice models, where in the Rimola's one¹³² more water molecules were implicated in the H^+ transfer-assistance and its surroundings, inferring stabilizing effects.

All these mentioned works based the $\text{NH}_2\text{CH}_2\text{OH}$ formation reaction on a concerted mechanism, in which the nucleophilic attack and the proton-transfer occurred synchronically. However, Chen and

1
2
3 Woon found that when the H₂CO and NH₃ reactants were well-encaged, fully surrounded by water
4 ice molecules, the C–N coupling took place spontaneously forming the NH₃⁺–CH₂O[–] zwitterionic
5 compound (see Figure 7C).¹³³ Zwitterions are neutral species bearing localized charges which are
6 stabilized by water solvent effects. In this case, 4 H₂O molecules are enough to induce the
7 barrierless formation of NH₃⁺–CH₂O[–] and its stabilization. Subsequent proton-transfer (assisted by
8 the water molecules) leading to final NH₂CH₂OH was computed to have an energy barrier of 13 kJ
9 mol^{–1} (B3LYP level and PCM). Similar zwitterion spontaneous formation was also observed more
10 recently by Riffet *et al.*,⁹³ where the H₂CO/NH₃/(H₂O)_n (n = 0–4) complexes were studied within the
11 PCM model at G3B3 DFT theory level. In this work, with only 3 H₂O molecules NH₃⁺–CH₂O[–]
12 formation was already observed. Calculated energy barrier of the water-assisted H⁺-transfer was 20
13 kJ mol^{–1} (for n = 4), 7 kJ mol^{–1} higher than that calculated by Chen and Woon.¹³³ Differences arise
14 from both the level of theory and the configuration of the water clusters, in which the Riffet's ones
15 resemble more to a water ice surface (see Figure 7D). For the sake of completeness, Riffet *et al.*
16 also investigated the formation of NH₃⁺–CH₂OH (protonated aminomethanol) by reaction of H₂CO
17 with the ammonium cation (NH₄⁺).⁹³ Computed barriers were found to be significantly higher
18 compared with the neutral processes, *e.g.*, 125 kJ mol^{–1} for n = 4.

19
20
21
22
23
24
25
26
27
28
29
30
31
32
33
34
35
36
37
38
39
40
41 As mentioned above, NH₂CH₂OH formation in water ice media can have a competitive channel,
42 *i.e.*, HOCH₂OH formation. This synthetic route was studied by Duvernay *et al.* on a water ice
43 surface model of 33 H₂O molecules at B3LYP level.¹³⁴ Results indicated that pure reaction of
44 H₂CO + H₂O → HOCH₂OH (in which the reacting H₂O belonged to the ice) occurred *via* formation
45 of a transient H₃O⁺/OH[–] pair, in which OH[–] performed the nucleophilic attack. This process was
46 calculated to have an energy barrier of ≈70 kJ mol^{–1}. However, the presence of NH₃ in the ice
47 catalyzed the reaction because of the easier formation of the NH₄⁺/OH[–] ion pair as intermediate
48 (energy barrier ≈8 kJ mol^{–1}), followed by the OH[–] nucleophilic attack (energy barrier ≈20 kJ mol^{–1})
49 forming deprotonated HOCH₂O[–] methyleneglycol. Protonation of HOCH₂O[–] was carried out by the
50
51
52
53
54
55
56
57
58
59
60

1
2
3 NH_4^+ (energy barrier $\approx 39 \text{ kJ mol}^{-1}$, the highest one). By comparing the energetics for $\text{NH}_2\text{CH}_2\text{OH}$
4
5 and HOCH_2OH formations, it is clear that the former process is more favorable. However, one has
6
7 to keep in mind that for $\text{NH}_2\text{CH}_2\text{OH}$ formation, NH_3 has to be in stoichiometric quantities with
8
9 H_2CO (it is the reactant), while for HOCH_2OH traces are enough to act as catalyst. Accordingly,
10
11 occurrence of one reaction or the other will strongly depend on the initial amount of NH_3 .
12
13

14
15 Formation of aminoalcohols from acetaldehyde (CH_3CHO) and acetone [$(\text{CH}_3)_2\text{CO}$] proceeds in a
16
17 similar way as for aminomethanol, *i.e.*, CH_3CHO and $(\text{CH}_3)_2\text{CO}$ react with NH_3 to give
18
19 $\text{NH}_2\text{CH}(\text{CH}_3)\text{OH}$ (1-aminoethanol) and $\text{NH}_2\text{C}(\text{CH}_3)_2\text{OH}$ (2-amino-2-propanol), respectively,
20
21 following the same “nucleophilic attack + proton-transfer” mechanism. These two reactions were
22
23 simulated by Fresneau *et al.* on different amorphous water-dominated dirty ice mantles at B3LYP-
24
25 D3 level,^{135,136} and by Chen & Woon¹³³ in an already mentioned work (see above). For all cases,
26
27 water acting as H^+ -transfer assistant was found to be essential to lower the energy barriers
28
29 compared with the gas phase processes. In the Fresneau’s works, both processes adopted a stepwise
30
31 mechanism, in which the first step involved the N–C coupling forming the $\text{NH}_3^+-\text{CH}(\text{CH}_3)\text{O}^-$ and
32
33 $\text{NH}_3^+-\text{C}(\text{CH}_3)_2\text{O}^-$ zwitterions followed by H^+ -transfer. As for the $\text{NH}_2\text{CH}_2\text{OH}$ case, in the Chen &
34
35 Woon work, zwitterion formations were spontaneous, at variance with respect to the Fresneau’s
36
37 ones. Differences can be explained by the different H_2O ice models: in the Fresneau’s model, both
38
39 the reactive species interacting with water and the same water molecules of the cluster were largely
40
41 geometrically constrained due to the H-bond network (hence representing more realistic ice surface
42
43 properties), while in the Chen and Woon’s ones, reactants were fully surrounded by H_2O molecules
44
45 and/or the clusters were geometrically exceedingly flexible, hence overestimating the stability of
46
47 the zwitterion induced by water. The highest calculated energy barriers for $\text{NH}_2\text{CH}(\text{CH}_3)\text{OH}$ and
48
49 $\text{NH}_2\text{C}(\text{CH}_3)_2\text{OH}$ formations were ≈ 34 and $\approx 26 \text{ kJ mol}^{-1}$, respectively, in the Fresneau’s works and
50
51 about $12\text{--}13 \text{ kJ mol}^{-1}$ in the work by Chen & Woon,¹³³ the energetic differences being due to the
52
53 different ice models (as explained above).
54
55
56
57
58
59
60

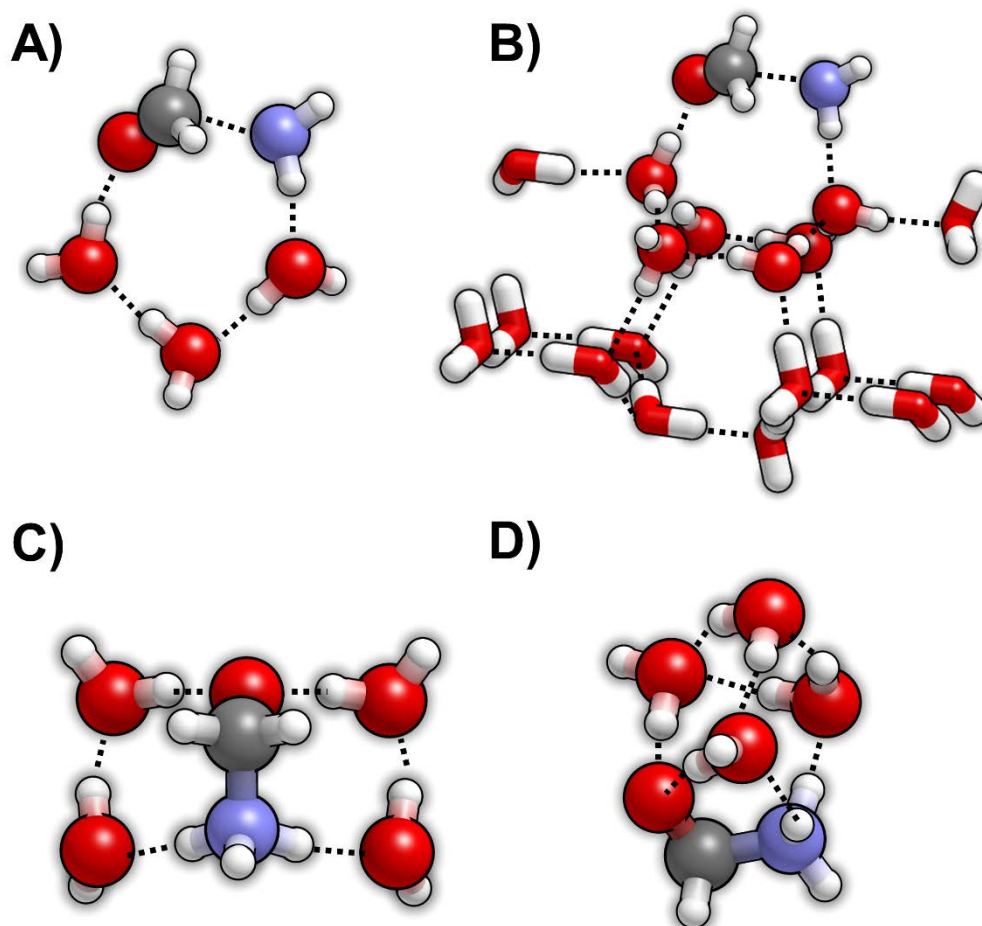


Figure 7. Representative initial structures for the formation of $\text{NH}_2\text{CH}_2\text{OH}$ from reaction of H_2CO with NH_3 : A) with a $(\text{H}_2\text{O})_3$ cluster model;¹³¹ B) with a $(\text{H}_2\text{O})_{18}$ cluster model as water ice surface (atoms involved in the H^+ -transfer assistance were highlighted as balls);¹³² C) with a $(\text{H}_2\text{O})_4$ cluster model;¹³³ and D) with a $(\text{H}_2\text{O})_4$ cluster model.⁹³ In these two later systems, the $\text{NH}_3^+-\text{CH}_2\text{O}^-$ zwitterion was spontaneously formed. Colour legend: oxygen in red, carbon in grey, nitrogen in blue and hydrogen in white.

3.5 Methanimine ($\text{CH}_2=\text{NH}$) formation and reactivity

The aforementioned works of Rimola *et al.*¹³² and Riffet *et al.*⁹³ also dealt with the dehydration of aminomethanol to form methanimine ($\text{CH}_2=\text{NH}$). In the former work,¹³² the $-\text{OH}$ and $-\text{NH}_2$ groups of aminomethanol acted as H-bond acceptor and donor groups, respectively, thus enabling the dehydration reaction through a water-assisted H^+ -transfer mechanism. Despite the catalytic behavior of water, the reaction presented an energy barrier of $\approx 90 \text{ kJ mol}^{-1}$. In the second work,⁹³ the same reaction was also studied using diverse water cluster models, presenting energy barriers of $100\text{-}150 \text{ kJ mol}^{-1}$. In this work, moreover, a charged mechanism was also studied involving the

1
2
3 dehydration of $\text{HOCH}_2\text{NH}_3^+$ (previously formed by reaction of NH_4^+ with H_2CO). This path
4
5 presented relatively lower energy barriers, between 70-90 kJ mol^{-1} .
6
7

8 Hydrogenation of HCN to form $\text{CH}_2=\text{NH}$ was studied by Woon following the reactions shown in
9
10 Scheme 4.¹¹⁹ The PESs of each step were characterised in absence of explicit water molecules but
11
12 using the PCM model at QCISD(T) level. Results indicated that for the first $\text{H}\cdot$ -addition, $\text{CH}_2\text{N}\cdot$
13
14 formation was more favorable than $\text{HCNH}\cdot$ (energy barriers of 30.5 vs 53.6 kJ mol^{-1} , respectively),
15
16 while the second one was considered to be barrierless.
17
18
19



20
21
22 **Scheme 4.** Hydrogenation of HCN leading to the formation of methanimine (in red). Compounds in
23
24 blue are the two possible radical intermediates.
25
26
27

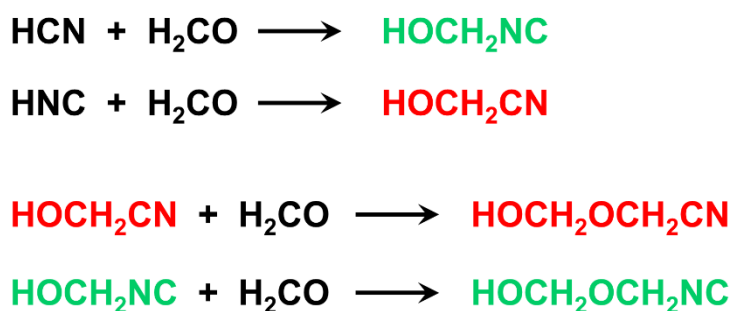
28
29 In relation to $\text{CH}_2=\text{NH}$ reactivity, its hydration (leading to $\text{NH}_2\text{CH}_2\text{OH}$) by reaction with one water
30
31 molecule and a second one acting as H^+ -assistant was calculated to have an energy barrier of 119
32
33 kJ mol^{-1} in PCM conditions, while its amination to give diaminomethane ($\text{CH}_2(\text{NH}_2)_2$) catalysed by
34
35 one H_2O molecule in PCM had an energy barrier of 83 kJ mol^{-1} .⁹¹ Reactivity of $\text{CH}_2=\text{NH}$ with
36
37 HCN giving aminoacetonitrile ($\text{NH}_2\text{CH}_2\text{CN}$) has also been studied. Since acetonitriles are a family
38
39 of iCOMs extensively investigated theoretically, formation of these compounds is reviewed in a
40
41 new section presented below.
42
43
44
45
46
47
48
49
50
51
52
53
54
55
56
57
58
59
60

3.6 Acetonitrile-derivatives formation

Acetonitrile is the compound with chemical formula CH_3CN . The most relevant feature is the nitrile $\text{C}\equiv\text{N}$ group. In this section, theoretical studies related to the formation of acetonitrile-derivatives on interstellar ice mantles are reviewed.

3.6.1 Hydroxyacetonitrile (HOCH_2CN) and hydroxyacetoinitrile (HOCH_2NC) formation

HOCH_2CN and HOCH_2NC are two hydroxylated acetonitrile compounds that can be formed by reaction of H_2CO with HCN and HNC , respectively. Although the reactants have widely been detected in the ISM,^{125,126,137} only HOCH_2CN has been detected very recently.¹³⁸ Woon examined these two reactions, as well as the subsequent reactivity of these compounds with a second H_2CO molecule (see Scheme 5), all of them in the presence of $(\text{H}_2\text{O})_n$ ($n = 0-2$) clusters at the MP2 level in PCM.¹³⁹ In the same work, the $\text{HCN} + \text{H}_2\text{O} \rightarrow \text{NH}=\text{CHOH}$ and $\text{HCN} + \text{NH}_3 \rightarrow \text{NH}=\text{CHNH}_2$ reactions were also investigated. Results indicated that reactivity with HNC presented lower energy barriers than with HCN (*e.g.*, 66 and 109 kJ mol^{-1} for the formation of HOCH_2CN from HNC and of HOCH_2NC from HCN on $(\text{H}_2\text{O})_2$ in PCM, respectively, see Scheme 5). Moreover, reaction energies were found to be negative in the former case, while slightly positive in the latter one. However, as shown above, HNC is more unstable than HCN (about 50 kJ mol^{-1}) and hence the difference in the energetic features of these reactions.



Scheme 5. Reactions simulated by Woon in the presence of small H_2O clusters.¹³⁹ Nitrile species are in red, while isonitriles in green.

1
2
3 More recently, Rimola and coworkers studied the HOCH₂CN formation by reaction of HCN with
4 H₂CO on a water ice model of 33 H₂O molecules at B3LYP-D3 level.¹⁴⁰ Results indicated that this
5 reaction was activated by a proton-transfer of the HCN to the H₂O ice, forming the H₃O⁺/CN⁻ ion
6 pair, in which CN⁻ was the responsible of the nucleophilic attack to the C atom of the C=O group
7 (see Figure 8). The resulting intermediate was ⁻OCH₂CN (deprotonated hydroxyacetonitrile), the
8 protonation of which was performed by recovering the proton initially transferred to the ice. Here,
9 the role of the ice was not as H⁺-transfer assistant but allowing the generation of the CN⁻ anion. The
10 first step presented the highest energy barrier, ≈54 kJ mol⁻¹, which is significantly lower to that
11 calculated by Woon.¹³⁹ Interestingly, authors also explored the formation of HOCH₂OCH₂CN by
12 reaction of the ⁻OCH₂CN intermediate (formed in the second step) with a second H₂CO molecule.
13 The coupling between the C atom of H₂CO with the charged O atom of ⁻OCH₂CN presented a very
14 low energy barrier (2.5 kJ mol⁻¹), indicating the feasibility of the process.
15
16
17
18
19
20
21
22
23
24
25
26
27
28
29
30

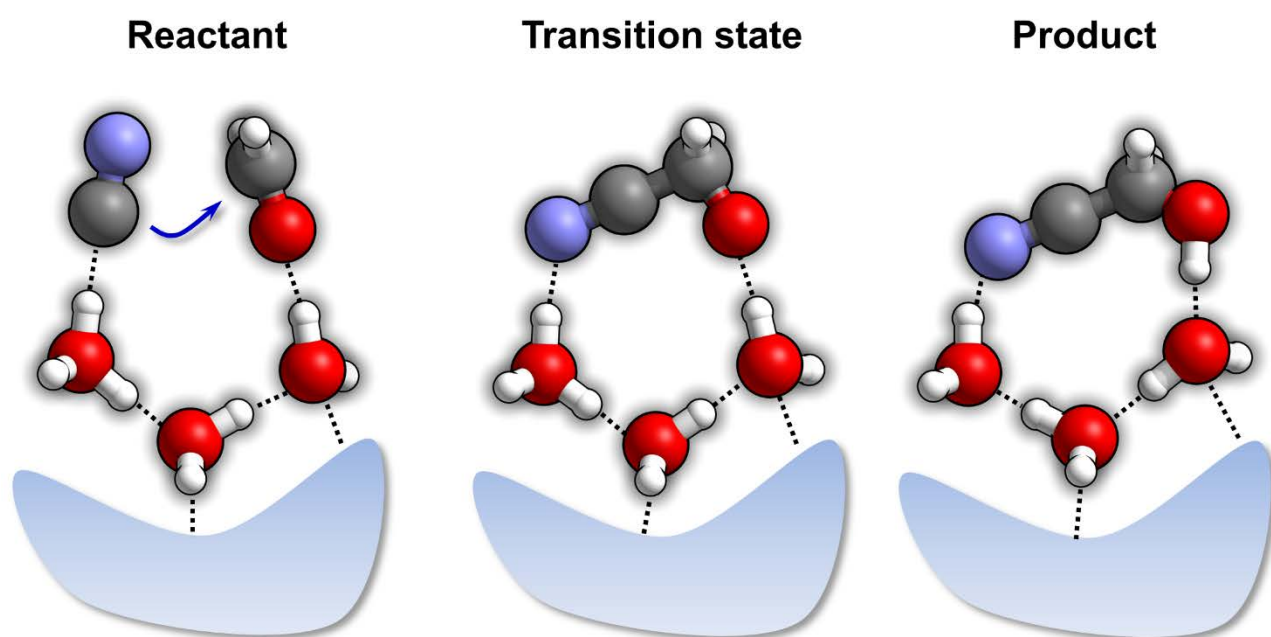


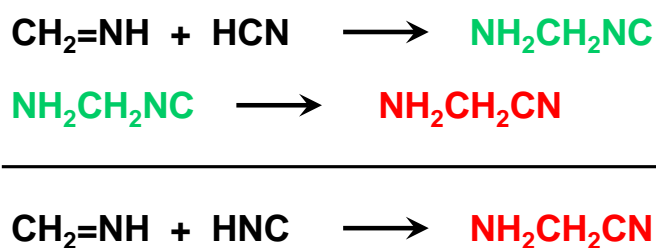
Figure 8. Schematic representation of the HOCH₂CN formation on a water ice surface through previous formation of the H₃O⁺/CN⁻ ion pair. The blue arrow indicates the nucleophilic attack of the C atom of CN⁻ anion to carbonyl group of H₂CO. Adapted from Ref. 140.

Rimola and coworkers also investigated the reactivity of HCN with CH₃CHO and (CH₃)₂CO leading to the formation of HOCH(CH₃)CN and HOC(CH₃)₂CN, respectively, in these cases

adopting diverse dirty ice surface model clusters.^{135,141} Simulated reaction mechanisms were the same as for the H₂CO case, presenting similar energy barriers (50-54 kJ mol⁻¹). Interestingly, calculations also indicated that the presence of traces of NH₃ in the ice favoured the formation of the CN⁻ anion because of the larger stability of the NH₄⁺/CN⁻ pair than the H₃O⁺/CN⁻ one, which was reflected by a significant decrease of the energy barriers (about 30-35 kJ mol⁻¹).

3.6.2 Aminoacetonitrile (NH₂CH₂CN) and aminoacetoisonitrile (NH₂CH₂NC)

NH₂CH₂CN and NH₂CH₂NC can be formed by reaction of CH₂=NH with HCN and HNC, respectively. Among these two nitriles, only NH₂CH₂CN has been detected in Sgr B2(N).¹⁴² Computational formation of these two compounds was first investigated by Woon at MP2 level in PCM conditions.⁹¹ The reaction mechanisms adopted are summarized in Scheme 6, in which NH₂CH₂CN formed *via* a two-step process, *i.e.*, reaction of CH₂=NH with HCN to give NH₂CH₂NC which then isomerized into NH₂CH₂CN. Nevertheless, direct formation of NH₂CH₂CN can occur by the addition of HNC to CH₂=NH. The two-step mechanism was computed to have energy barriers of 96 and 110 kJ mol⁻¹, while the direct one 40.5 kJ mol⁻¹.



Scheme 6. Formation of NH₂CH₂CN from CH₂=NH studied by Woon.⁹¹ Two-step mechanism (first addition of HCN to CH₂=NH to give NH₂CH₂NC and subsequent isomerization to NH₂CH₂CN, top) vs direct reaction with HNC (bottom). Nitriles in red, isonitriles in green.

In a more recent work, Koch *et al.* examined the NH₂CH₂CN formation by reaction of CH₂=NH with HCN in the gas phase, in the presence of (H₂O)_{*n*} (*n* = 1-3) clusters, and in the presence of 2 water molecules acting as proton transfer assistants plus 12 water molecules acting as ice spectators, all of them at B3LYP level.⁹² Authors considered two different mechanisms: a direct one, in which

1
2
3 $\text{NH}_2\text{CH}_2\text{CN}$ was formed by reaction of $\text{CH}_2=\text{NH}$ with HNC (previous $\text{HCN} \rightarrow \text{HNC}$
4 isomerization), and an indirect one, in which $\text{NH}_2\text{CH}_2\text{NC}$ was first formed, which then isomerized
5 into $\text{NH}_2\text{CH}_2\text{CN}$. The free energy profiles at 50 K for the gas phase reactions, in the presence of a
6 $(\text{H}_2\text{O})_2$ cluster, and with the “2+12” H_2O ice model are shown in Figure 9A and B for the direct and
7 indirect mechanisms, respectively, in which the gas phase optimized geometries are also shown.
8 The catalytic role of water is evident from these energy profiles, since energy barriers decrease
9 successively when the water ice model is improved. The indirect path was found to be more
10 favourable, with free energy barriers of 26 kJ mol^{-1} , while the direct one presented a free energy
11 barrier of 46 kJ mol^{-1} due to the initial $\text{HCN} \rightarrow \text{HNC}$ isomerization.
12
13
14
15
16
17
18
19
20
21
22
23
24
25
26
27
28
29
30
31
32
33
34
35
36
37
38
39
40
41
42
43
44
45
46
47
48
49
50
51
52
53
54
55
56
57
58
59
60

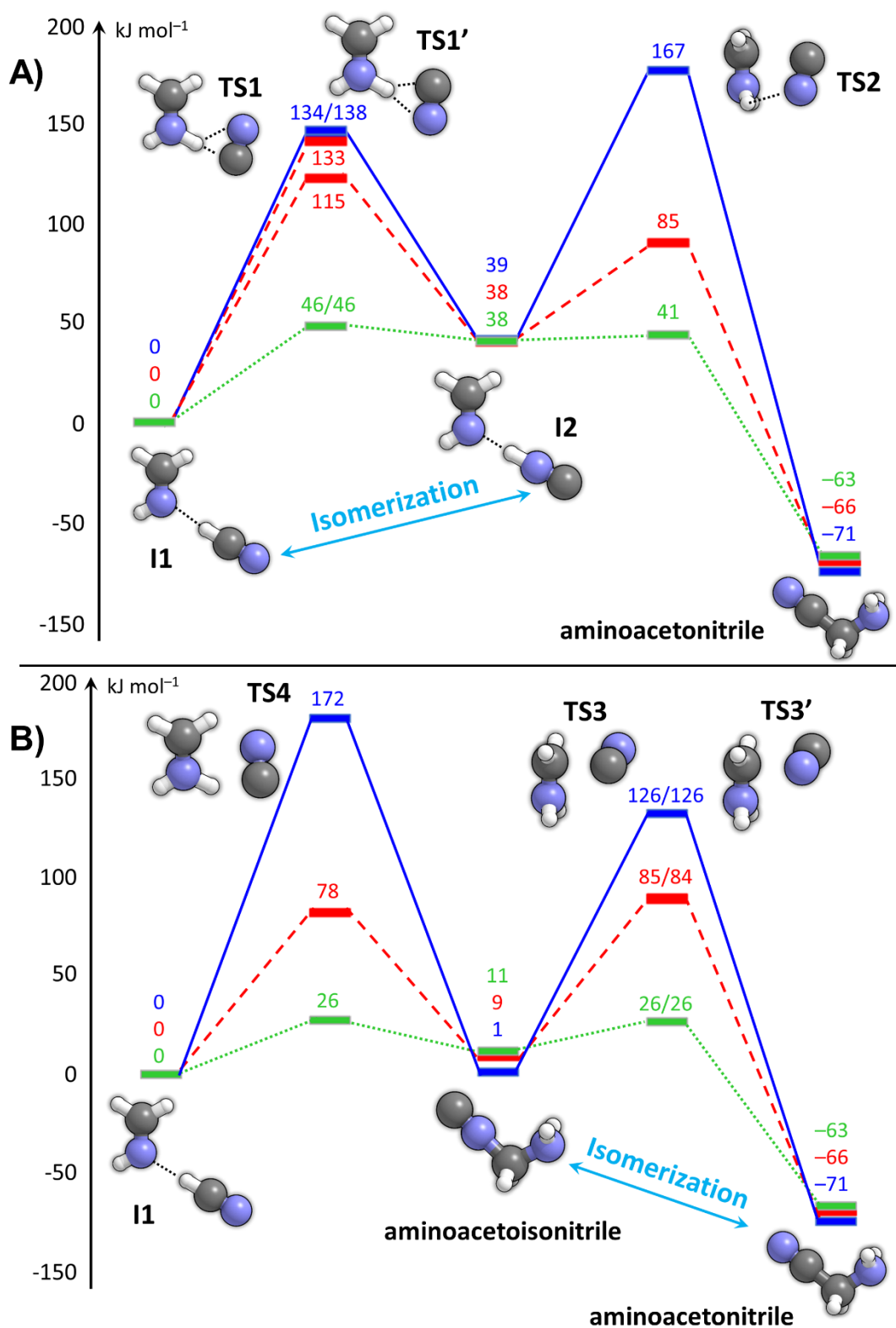
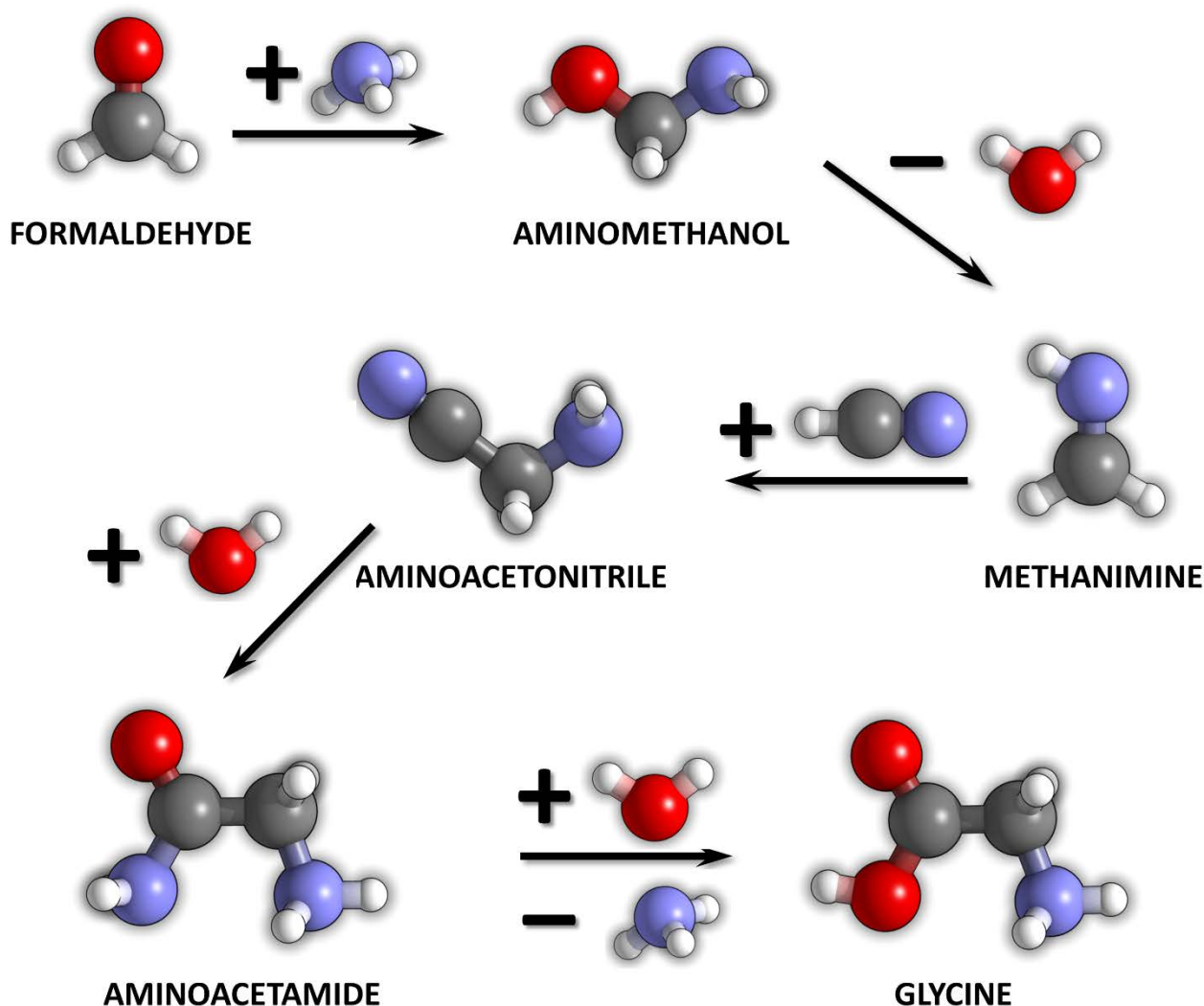


Figure 9. Free energy profiles (in kJ mol^{-1}) at 50 K for the direct (A) and indirect (B) formation of $\text{NH}_2\text{CH}_2\text{CN}$ from $\text{CH}_2=\text{NH}$ and HCN in gas phase (solid blue lines) in the presence of a $(\text{H}_2\text{O})_2$ cluster (dashed red lines) and in the presence of the “2+12” H_2O ice model (dotted green lines). Adapted from the work of Koch *et al.*⁹² We keep the same nomenclature of the stationary points of the original work. Optimized gas phase geometries are also shown. Atom color legends: carbon in grey, nitrogen in blue and hydrogen in white.

1
2
3 Finally, Rimola *et al.* also studied $\text{NH}_2\text{CH}_2\text{CN}$ formation on a water ice cluster model of 18
4 molecules at B3LYP level.¹³² They first simulated the direct path, namely, $\text{HCN} \rightarrow \text{HNC}$ first and
5 then $\text{CH}_2=\text{NH} + \text{HNC} \rightarrow \text{NH}_2\text{CH}_2\text{CN}$, whose energy barriers (adopting water-assisted H^+ -transfer
6 mechanisms) were ≈ 75 and ≈ 82 kJ mol^{-1} , respectively. In addition, authors also investigated an
7 ionic path. It started first with a proton-transfer from HCN to $\text{CH}_2=\text{NH}$ (energy barrier of ≈ 73 kJ
8 mol^{-1}), forming the $\text{CN}^-/\text{CH}_2\text{NH}_2^+$ ion pair, stabilized by interaction with the water ice. Then the
9 CN^- anion performed a nucleophilic attack to the C atom of CH_2NH_2^+ yielding final $\text{NH}_2\text{CH}_2\text{CN}$
10 with an energy barrier of 3 kJ mol^{-1} .
11
12
13
14
15
16
17
18
19
20
21

22 3.7 Glycine formation

23
24
25 Glycine ($\text{NH}_2\text{CH}_2\text{COOH}$), the simplest amino acid, has been identified in comets^{143–145} and its
26 presence in meteorites (among other amino acids) is usual.¹⁴⁶ The traditional route for the synthesis
27 of amino acids is the Strecker synthesis.¹⁴⁷ It involves different steps, some of them already
28 commented above: *i*) reaction of an aldehyde (RCHO , with R being a lateral chain) with ammonia
29 to give the corresponding aminoalcohol, *i.e.*, $\text{RCHO} + \text{NH}_3 \rightarrow \text{NH}_2\text{CH}(\text{R})\text{OH}$; *ii*) dehydration of the
30 aminoalcohol to give the corresponding imine, *i.e.*, $\text{NH}_2\text{CH}(\text{R})\text{OH} \rightarrow \text{CH}(\text{R})=\text{NH} + \text{H}_2\text{O}$; *iii*)
31 reaction of the imine with HCN to give the corresponding aminonitrile, *i.e.*, $\text{CH}(\text{R})=\text{NH} + \text{HCN} \rightarrow$
32 $\text{NH}_2\text{CH}(\text{R})\text{CN}$, and *iv*) acidic hydrolysis of the nitrile group which is converted firstly into an amide
33 ($-\text{CONH}_2$) and finally to an acid ($-\text{COOH}$) group, *i.e.*, $\text{NH}_2\text{CH}(\text{R})\text{CN} + 2\text{H}_2\text{O} \rightarrow \text{NH}_2\text{CH}(\text{R})\text{COOH}$
34 $+ \text{NH}_3$. These steps are schematically represented in Figure 10 for the particular case of glycine
35 formation, in which the aldehyde is H_2CO ($\text{R}=\text{H}$).
36
37
38
39
40
41
42
43
44
45
46
47
48
49
50
51
52
53
54
55
56
57
58
59
60



37 **Figure 10** Different steps of the Strecker synthesis of glycine from formaldehyde. Colour code:
38 oxygen in red, carbon in black, nitrogen in blue and hydrogen in white.
39

40
41 Interestingly, all the Strecker initial species (H_2CO , NH_3 , HCN , H_2O) are relatively abundant
42 interstellar molecules and Rimola *et al.* simulated this synthesis on a water ice surface model of 18
43 H_2O molecules in PCM.¹³² The first three steps, namely, formation of aminomethanol, methanimine
44 and aminoacetonitrile, have been commented above (Sections 3.4, 3.5 and 3.6.2, respectively), thus,
45 here we focus on the final step, the hydrolysis of the aminoacetonitrile, which involves the
46 successive nucleophilic attack of two H_2O molecules on the C atom of the nitrile. Calculated energy
47 barriers were found to be the highest ones of the overall process (159 and 163 kJ mol^{-1}), a fact
48 which led the authors to conclude that the entire Strecker synthesis is unlikely at the cryogenic
49
50
51
52
53
54
55
56
57
58
59
60

1
2
3 temperatures, advocating for external energy inputs such as UV radiation and cosmic rays to be
4
5 overcome.
6

7
8 These findings stimulated the same authors to investigate an alternative route for glycine formation
9
10 accounting for these energy inputs.¹²⁰ In Section 3.3.1 we reported formation of COOH· on
11
12 processed water ice clusters (see Figure 6). That work was within the context of glycine formation,
13
14 in which the next step after COOH· formation was its reactivity with CH₂=NH. On the RN cluster,
15
16 such a coupling, leading to formation of the NHCH₂COOH· radical, was computed to have an
17
18 energy barrier of 50 kJ mol⁻¹. On the RC cluster, authors identified an almost barrierless proton-
19
20 transfer from the H₃O⁺ species to CH₂=NH, and the formed CH₂=NH₂⁺ cation coupled to COOH·
21
22 through a lower energy barrier of 26 kJ mol⁻¹, to form the NH₂CH₂COOH·⁺ radical cation. Due to
23
24 the enhanced acidity of the CH₂ group of this radical cation, authors simulated that one H atom
25
26 could be transferred to the ice (energy barrier of ≈30 kJ mol⁻¹) so that the NH₂CHCOOH· radical
27
28 was ready to react with other radicals, to form different amino acids (*e.g.*, H· or CH₃· to give
29
30 glycine or alanine, respectively).
31
32
33
34

35
36 Alternative paths beyond the Strecker synthesis have also been computed. In a couple of works,
37
38 Nhlabatsi *et al.* investigated formation of interstellar glycine adopting two different channels based
39
40 on the CH₂=NH reactivity: *i*) CH₂=NH + CO + H₂O → NH₂CH₂COOH,¹⁴⁸ and *ii*) CH₂=NH + CO₂
41
42 + H₂ → NH₂CH₂COOH.¹⁴⁹ For both reactions, authors found a concerted mechanism in which all
43
44 the components reacted synchronically (Figure 11). Despite the elegance of these mechanisms, the
45
46 energy barriers were found to be 172 kJ mol⁻¹ (142 kJ mol⁻¹ if assisted by an additional H₂O
47
48 molecule) and 303 kJ mol⁻¹. For these two reactions, authors also investigated a stepwise
49
50 mechanism initiated by formation of the C(OH)₂ carbene (via CO + H₂O and CO₂ + H₂ reactions,
51
52 respectively), which upon reaction with CH₂=NH led to glycine. Although this later step was found
53
54 to have a relatively low energy barrier (38 kJ mol⁻¹), the processes were hindered by the high
55
56 energy barriers of the carbene formation (270 and 300 kJ mol⁻¹, respectively).
57
58
59
60

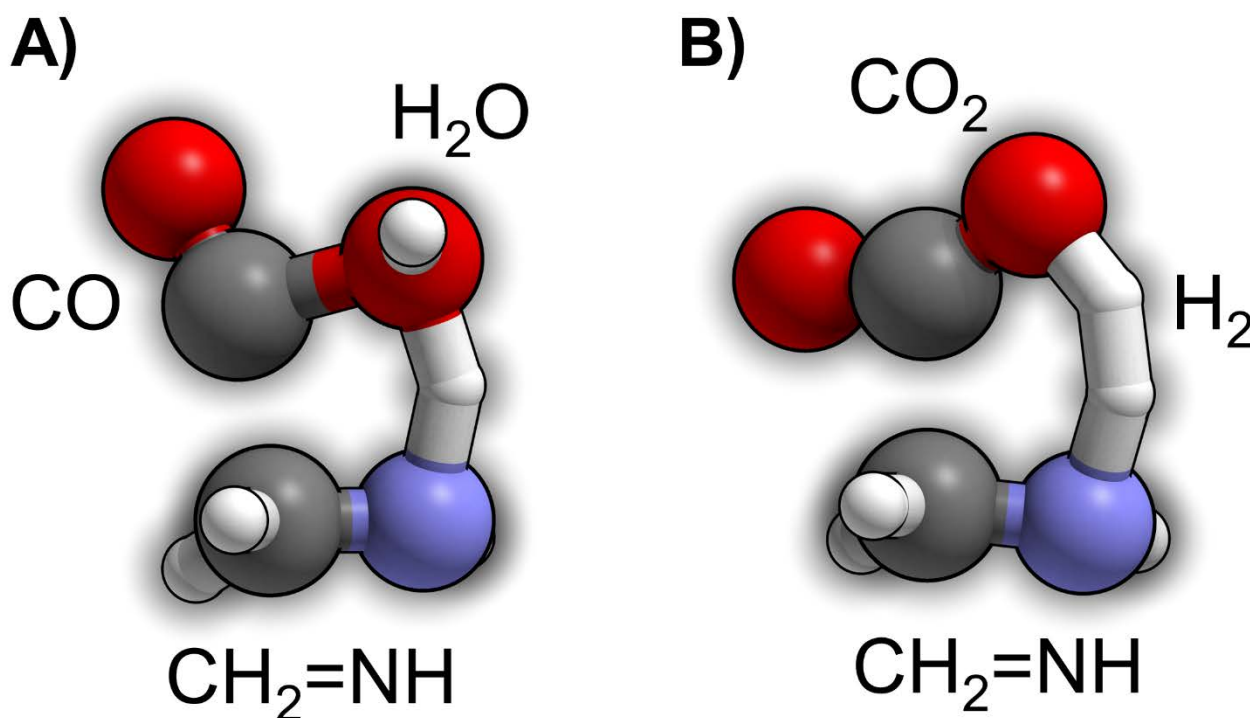
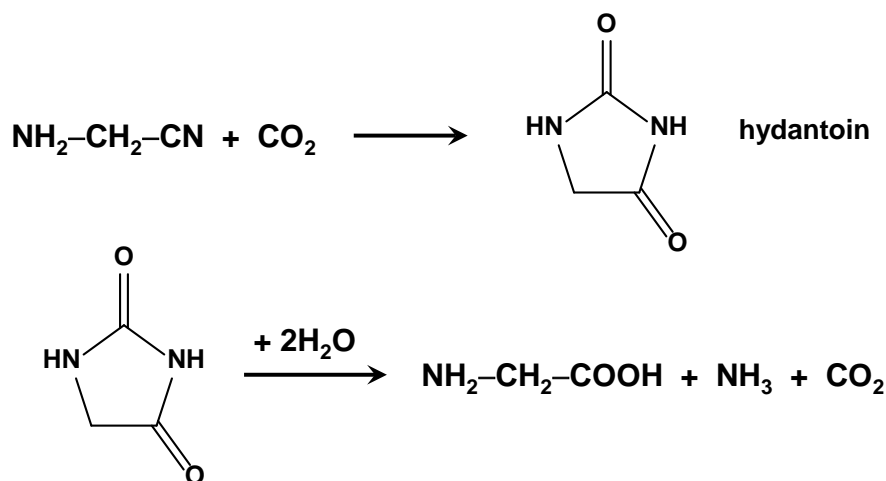


Figure 11 Transition states for the concerted formation of glycine from $\text{CH}_2=\text{NH} + \text{CO} + \text{H}_2\text{O}$ (A, Ref. 148) and $\text{CH}_2=\text{NH} + \text{CO}_2$ and H_2 (B, Ref. 149). Colour code: oxygen in red, carbon in black, nitrogen in blue and hydrogen in white.

Lee and Choe investigated the formation of glycine from HCN oligomers reacting with H_2O .¹⁵⁰ It was found that the HCN trimer, $\text{NH}_2\text{CH}(\text{CN})_2$, reacted with one H_2O molecule to form $\text{NH}_2\text{CH}(\text{CN})\text{CONH}_2$, and that water additions to this compound led to glycine formation through different paths, which were catalyzed by one water molecule assisting the H^+ -transfers. It was found the former reaction exhibited an energy barrier of 106 kJ mol^{-1} and that among the different investigated paths, the one involving $\text{NH}_2\text{CH}(\text{CN})\text{CONH}_2 + 2\text{H}_2\text{O} \rightarrow \text{glycine} + \text{HNCO} + \text{NH}_3$ was the most energetically favorable, with an overall energy barrier of 169 kJ mol^{-1} .

Finally, Kayanuma *et al.* investigated formation of glycine *via* a hydantoin mechanism.¹⁵¹ Hydantoin (2,4-imidazolidinedione) is an important precursor yielding glycine upon hydrolysis as it can be formed by CO_2 addition to aminoacetonitrile (see Scheme 7). Hydantoin has been identified in Murchison and Yamato-791198 meteorites.^{152,153} The reaction was simulated by the authors in the presence of two H_2O molecules acting as proton-transfer catalysts. Reactivity of aminoacetonitrile with CO_2 leading to hydantoin involved several steps, the highest energy barrier

being 111 kJ mol^{-1} . Hydantoin hydrolysis, performed by two H_2O molecules and accompanied by NH_3 and CO_2 elimination, exhibited intrinsic energy barriers between $176\text{-}255 \text{ kJ mol}^{-1}$. Authors pointed out that these energy barriers were too high to be overcome at cryogenic conditions even considering interstellar time-scales (10^6 years).



Scheme 7. Formation of hydantoin by reaction of aminoacetonitrile $\text{NH}_2\text{CH}_2\text{CN}$ with CO_2 (above) and its hydrolysis to give Gly (below).

Another interesting amino acid precursor is the hexamethylenetetramine (1,3,5,7-tetraazatricyclo-[3.3.1]-decane, $\text{C}_6\text{H}_{12}\text{N}_4$, HMT), a fourth-cycled molecule whose hydrolysis seem to form amino acids. Its solid-phase formation under astrophysical conditions has been simulated from reactivity of H_2CO and NH_3 in HCOOH -rich ices by Vinogradoff *et al.* in 2012.¹⁵⁴ The reaction involved first formation of $\text{NH}_2\text{CH}_2\text{OH}$ followed by water elimination to form $\text{CH}_2=\text{NH}$ and then reactivity between several $\text{CH}_2=\text{NH}$ molecules to form HMT. The identified mechanism consisted of a complex process with eight steps, in which the coexistence of $\text{CH}_2=\text{NH}$ and $\text{CH}_2=\text{NH}^+$ (this latter forming an ion pair with HCOO^-) was crucial to activate the HMT formation by $\text{CH}_2=\text{NH}$ additions. In addition to elucidating a plausible reactions mechanism, theoretical calculations also served to identify an intermediate species detected experimentally. The highest computed barrier is for the $\text{NH}_2\text{CH}_2\text{OH}$ dehydration to $\text{CH}_2=\text{NH}$, for which authors compute a thermal barrier at 70 K equal to 53.3 kJ mol^{-1} .

3.8 Nucleobases

Origin of nucleobases identified in meteorites is not well known. It seems, however, that formamide (NH_2CHO) is an essential precursor towards their formation. Several experimental works showed a selective reactivity of NH_2CHO on different mineral surfaces (*e.g.*, montmorillonite, titania, silica, etc.) towards formation of several nucleobases and derivatives.^{155–162}

The detailed formation mechanism of nucleobases is a matter of debate. A first paradigm postulates that NH_2CHO dehydrates first into HCN , the polymerization of which (or reaction with NH_2CHO) leads to nucleobase formation. Another one advocates that NH_2CHO polymerizes itself forming nucleobases and related species (see Figure 12).^{163–165}

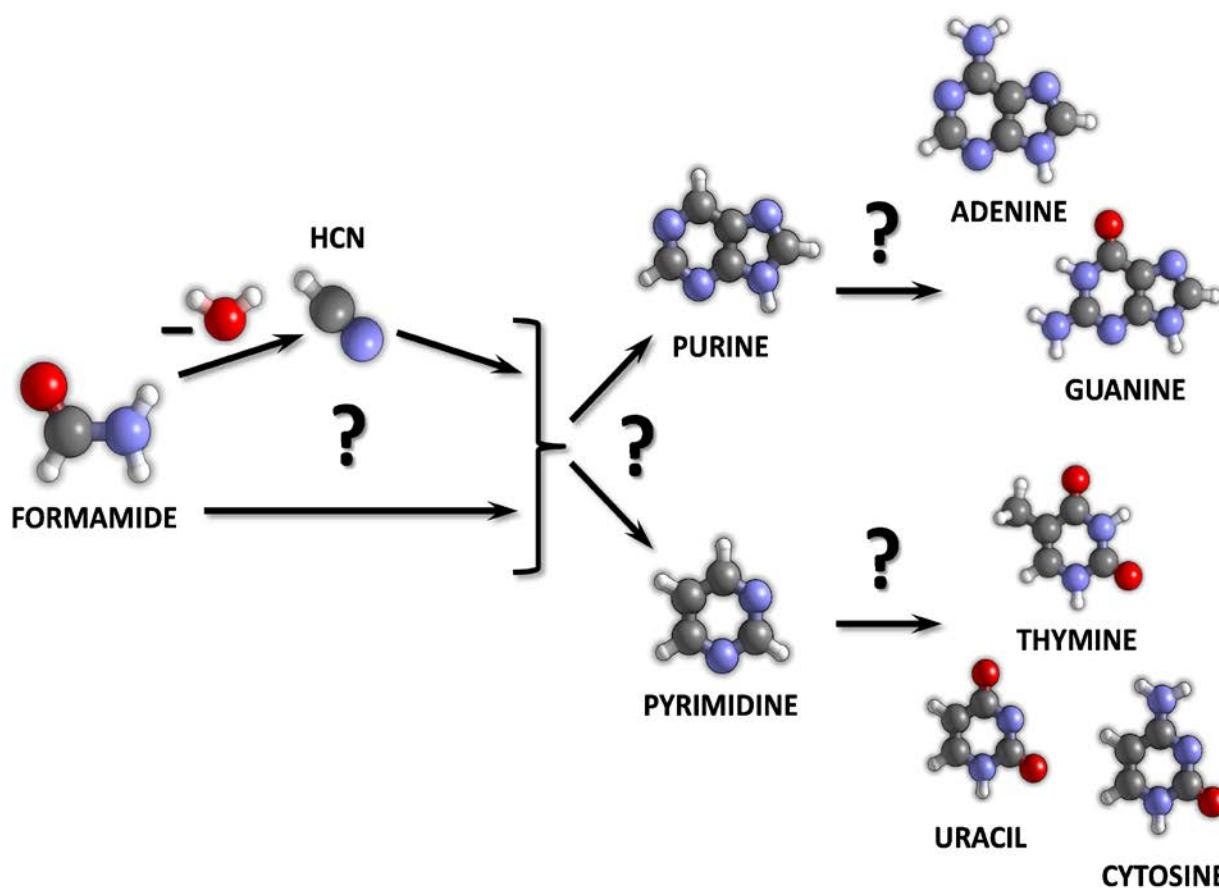
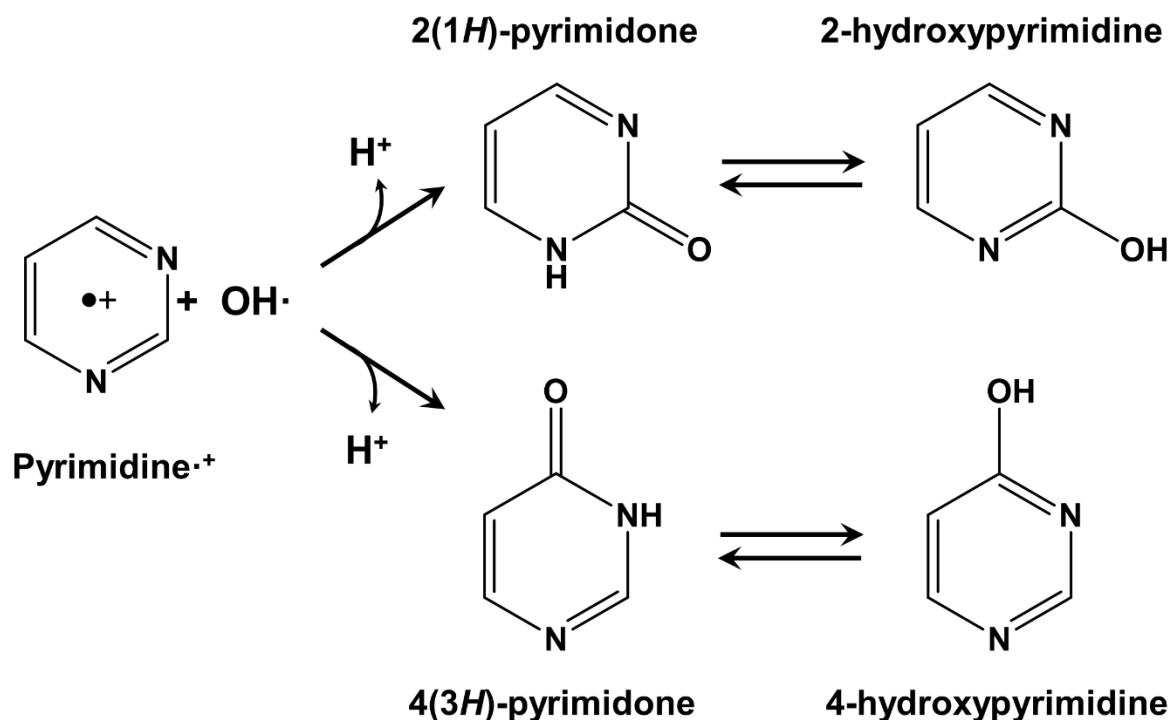


Figure 12. Pyrimidine, purine and nucleobases as products of either formamide decomposition into HCN followed by successive polymerization and reactivity or direct formamide polymerization and reactivity. Colour code: oxygen in red, carbon in black, nitrogen in violet and hydrogen in white.

1
2
3 Detailed mechanisms of nucleobase formation and the role of water in these reactions from an
4 atomistic point of view are very scarce. Nevertheless, few computational works can be found in
5
6 literature.
7
8

9
10 As regard the first paradigm, Ngueyn *et al.* have explored dehydration of formamide in presence of
11
12 1, 2 and 3 water molecules, at CCSD(T) level of theory.¹⁶³ Authors identified two possible routes: *i*)
13
14 one occurring through the H₂N–C–OH carbene species and forming HNC, and, *ii*) the other
15
16 occurring via the HN=CH–OH imine species and forming HCN. The latter is kinetically favored
17
18 since all the energy barriers are lower. In both routes, the role of the water was to favor the H⁺-
19
20 transfers, although barriers are too high in ISM conditions (the lowest one being 130 kJ mol⁻¹).
21
22 Despite this, in the case HCN to be formed, it was shown it could polymerize towards nucleobases
23
24 or related species.¹⁶³
25
26

27
28 Bera *et al.* investigated the formation of uracil (U) considering oxidation of pyrimidine (Py)
29
30 induced by UV photolysis at B3LYP and MP2 theory levels in absence and presence of a single
31
32 H₂O molecule.¹⁶⁶ Here, rather than providing full PESs, the thermodynamics of the different
33
34 mechanisms were investigated (*i.e.* no transition states and kinetic barriers were computed). The
35
36 first favourable step was the reaction of OH· (assumed to be derived from H₂O photolysis) with Py
37
38 or its radical cations form, Py·⁺ (also assumed to be formed by UV effects). Formation of different
39
40 mono-hydroxylated products were considered since OH· can react with any of the six positions of
41
42 Py/Py·⁺. The two most stable ones were 4(3*H*)-pyrimidone and 2(1*H*)-pyrimidone and their
43
44 tautomeric forms 4-hydroxypyrimidine and 2-hydroxypyrimidine (Scheme 8). The single water
45
46 molecule was found to help the H⁺ abstraction during the OH· addition. A second OH· addition into
47
48 these two stable products led to the formation of U as di-hydroxylated product (see Figure 12).
49
50 These reactions were also energetically favoured by the presence of water.¹⁶⁶
51
52
53
54
55
56
57
58
59
60

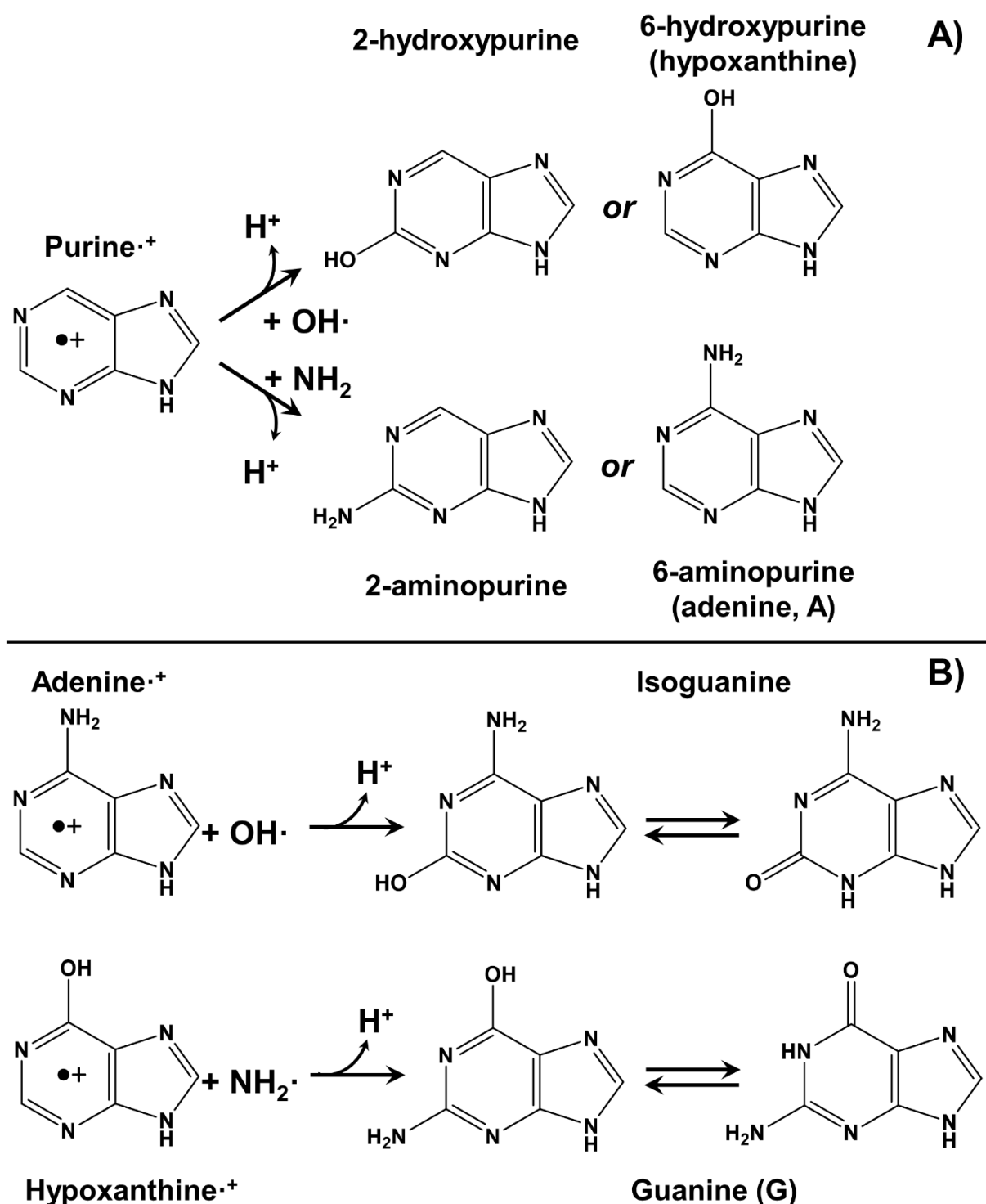


Scheme 8. Formation of mono-hydroxylated products from the $\text{OH}\cdot$ addition followed by H^+ elimination to the radical cation of pyrimidine. From Ref. 166.

In a more recent work, the same authors extended the calculations to study the formation of thymine (T, see Figure 12).¹⁶⁷ In this case, different radical routes were considered, combining two $\text{OH}\cdot$ additions with one $\text{CH}_3\cdot$ addition. It was found that the formation of T from Py through two successive $\text{OH}\cdot$ additions followed by a $\text{CH}_3\cdot$ addition was the most thermodynamically favourable path. The presence of an explicit H_2O molecule helped the H^+ abstraction of the intermediate species.

Formation of adenine (A) and guanine (G), from purine (Pu) was also explored by Bera *et al.* adopting the same idea to investigate the thermodynamic stability of different radical addition paths in the absence and presence of one H_2O molecule, in this case at B3LYP level and in PCM.¹⁶⁸ In particular, authors investigated the $\text{OH}\cdot$ and $\text{NH}_2\cdot$ additions onto the radical cation of Pu ($\text{Pu}\cdot^+$). The most stable products identified were 2-hydroxyhypurine and 6-hydroxypurine (hypoxanthine) for $\text{OH}\cdot$ addition, while 2-aminopurine and 6-aminopurine (adenine) for $\text{NH}_2\cdot$ addition (see Scheme 9). Interestingly, $\text{NH}_2\cdot$ addition onto hypoxanthine led to formation of G, while $\text{OH}\cdot$ addition to A

formed isoguanine. The role of the explicit water was again to help the H^+ abstractions, while PCM effects induce an additional stabilization of the products of these reactions.



Scheme 9. A) $OH\cdot$ and $NH_2\cdot$ additions followed by H^+ elimination to the radical cation of purine. One of the products is adenine. B) $OH\cdot$ and $NH_2\cdot$ addition followed by H^+ elimination to the radical cations of products formed in A). One of the products is guanine. From Ref. 168.

3.9 iCOMs from meteorite impacts

Cometary ices, similar to the interstellar ones, are predominated by H₂O but also contain other volatile species such as CO₂, NH₃ and CH₃OH.⁵¹ Additionally, recent analysis of dust samples from comet Wild 2 and 67/P showed the presence of glycine in the captured material^{145,169} These cometary ices could undergo high energy processes due to impacting with planetary surfaces. High energetic impacts generate strong pressure waves that propagate through the ice mantle, which eventually can ignite complex reactivity. This section reviews few computational works focused on the simulation of the chemistry taking place in an impacting cometary ice.

From a computational point of view, a high-pressure impact can be simulated by adopting multiscale shock-compression simulation technique, which is based on AIMDs in 3D periodic models. In this technique, periodic parameters are forced to shrink within a certain short time and then relaxed, *i.e.*, expanded and cooled down up to thermalization of the systems. Such a procedure was carried out by Goldman and coworkers to investigate the chemistry triggered by high-pressures of a mixture ice with composition of 20H₂O, 10CH₃OH, 10NH₃, 10CO and 10CO₂ *per* 3D unit cell.^{51,52}

In a former study,⁵¹ authors adopted expensive AIMDs introducing shock velocities of 5, 6, 7, 8, 9, 10 km s⁻¹ for 5–11 ps (2 ps only for 10 km s⁻¹), corresponding to initial impact pressures of 10, 18, 24, 37, 47 and 59 GPa, and temperatures of 706, 1196, 1590, 2549, 3141 and 4083 K, respectively. After the impact period, systems were decompressed and cooled down (relaxed) for some ps. Authors analyzed the eventual formation of new bonds among the initial reactant mixture. The most interesting finding of this work was the formation of a Gly-related species at 9 km s⁻¹ (47 GPa): the high-pressure impact caused the formation of a long C–N-bonded oligomer containing the –NH–CH₂–COOH sequence. In the relaxation phase, such oligomer broke apart forming several C–N-bearing species such as HCN and NH₂–COOH, but the sequence corresponding to Gly remained intact. This molecular complex could eventually react with protonated species to form glycine; for

1
2
3 instance, ${}^{-}\text{OCO-NH-CH}_2\text{-COOH} + \text{H}_3\text{O}^+ / \text{NH}_4^+ \rightarrow \text{NH}_2\text{-CH}_2\text{-COOH} + \text{CO}_2 + \text{H}_2\text{O/NH}_3$, with free
4
5
6 reaction energies of $-101.3/-9.2 \text{ kJ mol}^{-1}$.
7

8
9 An analogous procedure was adopted in a second work,⁵² but at Density Functional Tight Binding
10
11 (DFTB) level, which is a simplified version of standard QM calculations which is close to
12
13 molecular mechanics in terms of computer resources. With the same initial ice composition of the
14
15 previous work, shocks at 36, 48, 60 GPa for 100 ps (phase 1), followed by adiabatic expansions
16
17 (phase 2) and final cooling at 300 K (phase 3) were simulated. In phase 1, several new C-C and C-
18
19 N bond forming species were identified, the actual composition depending on the given pressure.
20
21 Some of these transient species decomposed during phase 2. Among survival species, in phase 3,
22
23 authors identified amino acids precursors as well as aliphatic and aromatic hydrocarbons.
24
25
26
27
28
29
30
31
32
33
34
35
36
37
38
39
40
41
42
43
44
45
46
47
48
49
50
51
52
53
54
55
56
57
58
59
60

Table 1 Summary of all the reviewed works. Acronyms legend: CO, carbon monoxide; CH₃OH, methanol; H₂CO formaldehyde; HCOOH, formic acid; CO₂⁻, carboxylate anion; CO₂, carbon dioxide; FoAm, formamide; HCO·, formyl radical; NH₂·, amino radical; CH₃· methyl radical; CN·, cyanil radical; AM, aminomethanol; MG, methylenglicole; POM, polyoxymethylene; AcAl, acetaldehyde; Ac, acetone; HCN, hydrogen cyanide; HNC, hydrogen isocyanide; Gly, glycine; NH₃, ammonia; NH₄⁺, ammonium cation; CH₂=NH, methanimine; CH₃NH₂, methylamine; HOCN, cyanic acid; HNCO, isocyanic acid; CH₃COOH, acetic acid; HAN, hydroxyacetonitrile; HAisoN, hydroxyacetoisonitrile; AAN, aminoacetonitrile; HMT, hexamethylenetetramine; Py, pyrimidine; Pu, purine; U, uracil; A, adenine; G, guanine; T, thymine.

Topic	Molecules	QM Method	Ice model	Reference
Hydrogenation of CO to form H ₂ CO and CH ₃ OH	CO, H ₂ CO, CH ₃ OH	MP2, QCISD, ONIOM	1-4, 12H ₂ O PCM	97
Hydrogenation of CO to form H ₂ CO and CH ₃ OH	CO, H ₂ CO, CH ₃ OH	B3LYP, CCSD(T)	3, 18, 32H ₂ O	98
Hydrogenation of CO to form H ₂ CO and CH ₃ OH	CO, H ₂ CO, CH ₃ OH	QM/MM	No water: on silica surfaces	99
Hydrogenation of CO to form H ₂ CO and CH ₃ OH	CO, H ₂ CO, CH ₃ OH	QM/MM	No water: on silica surfaces	100
Formation of CH ₃ OH, HCOOH and CO ₂ through cationic reactions	CH ₃ OH, FH, CO ₂	B3LYP, MP2	4, 17H ₂ O	101
Hydrogenation of HNCO to form FoAm	HNCO, FoAm	QM/MM	Hemispherical cluster cut from amorphous slab	64
Formation of FoAm from HCO· + NH ₂ ·, HCN + H ₂ O and CN· + H ₂ O	HCO·, NH ₂ ·, HCN, CN·, FoAm	BHLYP	33H ₂ O	114
Formation of AcAl from HCO· + CH ₃ ·	HCO·, CH ₃ ·, AcAl	M06-2X-D3 B3LYP	18, 33H ₂ O	117
Formation of FoAm from CO + 2NH ₃	CO, NH ₃ , FoAm	B3LYP	No water	118
Hydrogenation of HCN to form CH ₂ =NH Radical formation of HCOOH and CO ₂	HCN, CH ₂ =NH, HCOOH, CO ₂	QCISD, QCISD(T)	PCM	119
Reactions of FH and CH ₂ =NH with NH ₃ Direct formation of Gly from HCOOH + CH ₂ =NH	HCOOH, CH ₂ =NH, NH ₃ , Gly	MP2	1, 2H ₂ O PCM	91
Formation of NH ₄ ⁺ /CO ₂ ⁻ by protonation of NH ₃ from HCOOH	HCOOH, NH ₃ , NH ₄ ⁺ /CO ₂ ⁻	B3LYP	2-7, 9, 14, 15H ₂ O	121

Formation of a $\text{CH}_3\text{NH}_2^+/\text{CO}_2^-$ ion pair	CH_3NH_2 , CO_2 , $\text{CH}_3\text{NH}_2^+/\text{CO}_2^-$	B3LYP	$n\text{H}_2\text{O}$, $0 \leq n \leq 20$	122
Formation of $\text{NH}_4^+/\text{OCN}^-$ by protonation of NH_3 from HOCN/HNCO.	HOCN, HNCO, NH_3	B3LYP, ONIOM	2-15 H_2O PCM	123
Reproduction of "XCN" interstellar band	HOCN, HNCO, NH_3	B3LYP	2-12 H_2O	124
$\text{HCN} \leftrightarrow \text{HNC}$ isomerization	HCN, HNC	MP2, CCSD(T)	1-4 H_2O	127
$\text{HCN} \leftrightarrow \text{HNC}$ isomerization	HCN, HNC	B3LYP	Reaction core: 3 H_2O Solvation: 12 H_2O , PCM	94
Deprotonation of CH_3COOH , HCN, HNC in water ice	CH_3COOH , HCN, HNC	B3LYP, MP2	2-6 H_2O	128
Formation of AM from $\text{H}_2\text{CO} + \text{NH}_3$ Formation of MG from $\text{H}_2\text{CO} + \text{H}_2\text{O}$	H_2CO , AM, MG	MP2	PCM	129
Formation of POM from AM	AM, POM derivatives	MP2	PCM	130
Formation of AM from $\text{H}_2\text{CO} + \text{NH}_3$	H_2CO , AM	B3LYP, MP2, CCSD(T)	1-3 H_2O	131
Reactivity of H_2CO , AcAl and Ac with NH_3	H_2CO , AcAl, Ac, aminated products	B3LYP	2, 4, 9, 12 H_2O PCM	133
Two first steps of glycine Strecker's synthesis	H_2CO , AM, $\text{CH}_2=\text{NH}$,	B3LYP	1-4 H_2O	93
Formation of MG from $\text{H}_2\text{CO} + \text{H}_2\text{O}$ (with NH_3 presence)	H_2CO , MG	B3LYP	33 H_2O	134
Formation of $\text{H}_2\text{NCH}(\text{CH}_3)\text{OH}$ from AcAl + NH_3 Formation of $\text{NCCH}(\text{CH}_3)\text{OH}$ from AcAl + HCN	AcAl, HCN, NH_3	B3LYP-D3	12 H_2O	135
Formation of $\text{HOC}(\text{CH}_3)_2\text{NH}_2$ from Ac + NH_3	Ac, NH_3	B3LYP-D3	15 H_2O	136
Formation of HAN and HAisoN from $\text{H}_2\text{CO} + \text{HNC}/\text{HCN}$ and isomerization	H_2CO , HCN/HNC, HAN, HAisoN	MP2	PCM	139
Formation of HAN, HAisoN and POM-CN from $\text{H}_2\text{CO} + \text{HNC}/\text{HCN}$	H_2CO , HCN, POM-CN	B3LYP	33 H_2O	140

1 2 3 4	Formation of $\text{HOC}(\text{CH}_3)_2\text{CN}$ from Ac + HCN (with NH_3 presence)	Ac, HCN, NH_3	B3LYP-D3	$33\text{H}_2\text{O}$	141
5 6 7	Formation of AAN from $\text{CH}_2=\text{NH}$ + HCN/HNC	MI, HCN/HNC, AAN	B3LYP	Reaction core: $3\text{H}_2\text{O}$ Solvation: $7\text{H}_2\text{O}$, PCM	92
8 9	Glycine Strecker's synthesis	H_2CO , AM, $\text{CH}_2=\text{NH}$, AAN, Gly	B3LYP	$18\text{H}_2\text{O}$	132
10 11	Radical formation of Gly from HCOOH and MI	HCOOH, $\text{CH}_2=\text{NH}$, Gly	BHLYP	$8\text{H}_2\text{O}$	120
12 13 14	Formation of Gly from $\text{CH}_2=\text{NH}$ + CO + H_2O	$\text{CH}_2=\text{NH}$, CO	B3LYP and several others	$1-4\text{H}_2\text{O}$	148
15 16	Formation of Gly from $\text{CH}_2=\text{NH}$ + CO_2 + H_2	$\text{CH}_2=\text{NH}$, CO, CO_2 , H_2	B3LYP and several others	No water	149
17 18	Formation of Gly from $\text{H}_2\text{NCH}(\text{CN})\text{CONH}_2$	$\text{H}_2\text{NCH}(\text{CN})\text{CONH}_2$, Gly	B3LYP, CBS-QB3	$1\text{H}_2\text{O}$	150
19 20	Formation and Gly <i>via</i> hydantoin	AAN, hydantoin, Gly	B3LYP	$1-3\text{H}_2\text{O}$	151
21 22 23	Formation of HMT from reactivity of $\text{CH}_2=\text{NH}$	H_2CO , HMT, $\text{CH}_2=\text{NH}$, HCOOH, AM, NH_3	B3LYP	$\text{NH}_3\text{:H}_2\text{CO}\text{:FH}$ mixed clusters but no ice model	154
24 25	Dehydration of FoAm to form H_2O + HCN/HNC	FoAm, HCN/HNC	MP2, CCSD(T)	$1-3\text{H}_2\text{O}$	163
26 27 28	Interaction and dehydration of FoAm on silica surfaces. No kinetic barriers	FoAm	PBE	No water: silica surfaces	170
29 30	Radical formation of U from Py. No kinetic barriers	Py, U	B3LYP, MP2	No water	166
31 32	Radical formation of T from Py (<i>via</i> U). No kinetic barriers	Py, T, U	B3LYP, MP2	No water	167
33 34 35	Formation of A and G from Pu. No kinetic barriers	Pu, A, G	B3LYP	PCM	168
36 37	Formation of iCOMs in impacting cometary ices	H_2O , CH_3OH , NH_3 , CO, CO_2	B3LYP	$20\text{H}_2\text{O}\text{:}10\text{CH}_3\text{OH}\text{:}10\text{NH}_3\text{:}10\text{CO}\text{:}10\text{CO}_2$	51
38 39 40	Formation of iCOMs in impacting cometary ices	H_2O , CH_3OH , NH_3 , CO, CO_2	DFTB	$20\text{H}_2\text{O}\text{:}10\text{CH}_3\text{OH}\text{:}10\text{NH}_3\text{:}10\text{CO}\text{:}10\text{CO}_2$	52

1
2
3
4
5
6
7
8
9
10
11
12
13
14
15
16
17
18
19
20
21
22
23
24
25
26
27
28
29
30
31
32
33
34
35
36
37
38
39
40
41
42
43
44
45
46

4 Conclusions and future perspectives

In this work, most of the quantum mechanical studies addressing the formation of iCOMs on ice mantles have been reviewed. They are not only focused on standard iCOMs but also to simpler organic compounds as well as those of increased complexity, *i.e.*, formation of H₂CO and CH₃OH, NH₂CHO, acidic organic species (*e.g.*, HCOOH), aminoalcohols, CH₂=NH, acetonitriles, glycine and nucleobases. The different reaction-types yielding their formation have also been revised theoretically: hydrogenations, radical additions, radical-radical couplings, and ion-ion, ion-neutral and neutral-neutral reactions. All the reviewed works, including useful details, are summarized in Table 1.

Since water is the main constituent of interstellar ices, ice mantles were simulated by either explicit water molecules or implicitly with PCM solvation models. It was shown that water exerted from moderate to strong catalytic effects in the reactions. They were particularly important when water molecules were explicitly considered due to their role as H⁺-transfer assistants, in which the energy barriers decreased as a consequence of the lower geometrical strains in TS structures than in gas phase. In other cases, water stabilized ion pairs, allowing the occurrence of ion-induced reactions. Despite these catalytic effects, most of the energy barriers were calculated to be significantly high to occur at typical temperatures of MCs (10-20 K) and, accordingly, activation by temperature was in most of the cases claimed.

All the reviewed works have contributed to improve our know-how of the iCOMs formation on ice mantles, by figuring out the processes from a molecular standpoint, providing exclusive structural and energetic features, and helping us to assess their feasibility under interstellar conditions. However, several relevant aspects remain still missing.

One of them deals with the plausibility of the occurrence of water-assisted H⁺-transfer processes adopting a relay mechanism on ice surface mantles. To take place, the implicated waters must be connected by H-bond interactions in a suitable way as they can be capable to donate and receive H⁺

1
2
3 properly. However, whether this situation is indeed present or not in actual ice mantles and how the
4 structural state and the presence of other ice components can affect this water catalytic property are
5 still open questions that require further investigations, to be possible by combining experimental
6
7
8
9
10 measurements with quantum chemical calculations.

11
12 Another interesting aspect is the reliability of the surface models representing the ice mantles.
13 Among the reviewed works, they consisted of either minimal $(\text{H}_2\text{O})_n$ ($n = 1-4$) clusters, in which the
14 mobility of the H_2O molecules was at its maximum, larger clusters ($n > 18-20$), and only in the
15 most recent works they were represented by amorphous clusters of hundreds of H_2O molecules,
16
17
18
19
20
21 while adoption of periodic slab models is very scarce. However, as mentioned in the Computational
22 Framework section, theoretical results can dramatically depend on the ice model and the evaluation
23 of this aspect, *i.e.*, how results are actually affected by the structure and type of the ice model,
24
25
26
27
28
29 should represent an important topic of future works dealing with iCOMs formation on atomistic ice
30 models. Accordingly, comprehensive studies, and consistent from a methodological viewpoint,
31
32
33
34
35
36 assessing the reliability of the different ice models and analyzing how similar/different are the
37 results when using different ice models, are of great importance.

38
39 The role of water ice as catalyst has been clearly evidenced here. However, ice mantles are not
40 limited to this role only. For instance, they can also act as reactant suppliers. In the reviewed works
41
42
43
44
45
46
47
48
49
50
51
52
53
54
55
56
57
58
59
60 this role was shown when the reactants were also usual ice components, *e.g.*, H_2O itself or CO and
 NH_3 . However, there are two other roles which have hardly been investigated. One is as reactant
concentrator. Indeed, ice surfaces can immobilize and concentrate species, keeping them in close
proximity for subsequent reactions. Assessing this role can be carried out by calculating the
interaction energies between the reactants and the ice surfaces, which can indicate how strongly
reactants are retained on the surfaces. Interaction energies were usually provided in most of the
reviewed works but their relationship with the capability of surface ices to act as reactant suppliers
is not usual. Another way to assess this role is by simulating the diffusivity of the reactive species.

1
2
3 This can be performed with AIMDs, in which retention times can be provided. Nevertheless,
4
5 AIMD-based studies devoted to the diffusion properties of the reactants are very scarce. The other
6
7 role is as third body, *i.e.*, ices quickly absorb the reaction energy excess, thereby stabilizing the
8
9 product. This role can be investigated theoretically with AIMDs at NVE, in which the total energy
10
11 E is conserved along the whole simulation. These simulations allow elucidating how the nascent
12
13 reaction energies are partitioned, *i.e.*, what amount transforms into translational and internal
14
15 energies of the product and what dissipates among the ice. Studies focused on this aspect are also
16
17 very rare. Remarkably, the lack of this kind of works also evidences that use of AIMDs is very
18
19 scarce in iCOMs formation investigations, a critical aspect since dynamic effects can be of great
20
21 relevance especially in those reactions in which thermal heating is essential.
22
23
24

25
26 Finally, we address some words claiming for the need to simulate non-investigated reactions.
27
28 Several important synthetic routes have indeed been simulated successfully but others, which are
29
30 also important, are still missing. For instance, radical-radical couplings have scarcely been
31
32 investigated: the $\text{HCO}\cdot + \text{NH}_2\cdot$ and $\text{HCO}\cdot + \text{CH}_3\cdot$ reactions have been simulated,^{114,117} while other
33
34 radical couplings are still to be studied. This is quite surprising since these reactions are assumed to
35
36 be the main channels to form iCOMs usually detected in diverse astrophysical objects, as mentioned
37
38 in the Introduction. In the same line, computational simulations have also been useful to identify
39
40 new formation paths which are not normally accounted for in astrochemical modelling schemes
41
42 (*e.g.*, $\text{CN}\cdot + \text{H}_2\text{O}$,¹¹⁴). Moreover, for some identified iCOMs, no reaction mechanisms have been
43
44 proposed and simulated (*e.g.*, acetone $(\text{CH}_3)_2\text{CO}$ or vinyl alcohol $\text{CH}_2=\text{CHOH}$). Because of that,
45
46 extensive quantum mechanical simulations devoted to novel “on-surface” formation paths to check
47
48 their plausibility will be of great value. In relation to cometary and meteoritic biomolecules, the
49
50 focus has been done essentially on glycine formation (while no simulations have been done for the
51
52 rest of amino acids), on particular paths for nucleobases (but the NH_2CHO -based routes are
53
54 unexplored yet), whereas sugars formation routes have not been addressed. Moreover,
55
56
57
58
59
60

1
2
3 understanding the role of the cometary and meteoritic minerals and ices will help us to get a better
4
5 understanding on the origin of these compounds.
6
7

8 **5 Acknowledgements**

9
10 Albert Rimola is indebted to “Ramón y Cajal” program. This research was funded by MINECO
11 (project CTQ2017-89132P), AGAUR (project 2017SGR1320), MIUR (Ministero dell’Istruzione,
12 dell’Università e della Ricerca) and Scuola Normale Superiore (project PRIN 2015, STARS in the
13 CAOS - Simulation Tools for Astrochemical Reactivity and Spectroscopy in the
14 Cyberinfrastructure for Astrochemical Organic Species, cod. 2015F59J3R). This project has
15 received funding from the European Research Council (ERC) under the European Union's Horizon
16 2020 research and innovation programme, for the Project “The Dawn of Organic Chemistry”
17 (DOC), grant agreement No 741002.
18
19
20
21
22
23
24
25
26
27
28

29 **6 Author’s information**

30 **Corresponding author**

31
32
33 Albert Rimola. Email: albert.rimola@uab.cat
34
35
36
37

38 **Other authors**

39
40 Lorenzo Zamirri. Email: lorenzo.zamirri@unito.it
41
42

43 Piero Ugliengo. Email: piero.ugliengo@unito.it
44

45 Cecilia Ceccarelli. Email: cecilia.ceccarelli@univ-grenoble-alpes.fr
46
47

48 **ORCID**

49 Lorenzo Zamirri: 0000-0003-0219-6150
50

51 Piero Ugliengo: 0000-0001-8886-9832
52

53 Albert Rimola: 0000-0002-9637-4554
54

55 Cecilia Ceccarelli: 0000-0001-9664-6292
56
57
58
59
60

7 References

- (1) McGuire, B. A. 2018 Census of Interstellar, Circumstellar, Extragalactic, Protoplanetary Disk, and Exoplanetary Molecules. *Astrophys. J. Suppl. Ser.* **2018**, *239*, 17 (48 pp).
- (2) Herbst, E.; van Dishoeck, E. F. Complex Organic Interstellar Molecules. *Annu. Rev. Astron. Astrophys.* **2009**, *47*, 427–480.
- (3) Ceccarelli, C.; Caselli, P.; Fontani, F.; Neri, R.; López-Sepulcre, A.; Codella, C.; Feng, S.; Jiménez-Serra, I.; Lefloch, B.; Pineda, J. E.; et al. Seeds Of Life In Space (SOLIS): The Organic Composition Diversity at 300–1000 Au Scale in Solar-Type Star-Forming Regions. *Astrophys. J.* **2017**, *850*, 176 (15 pp).
- (4) Ceccarelli, C.; Loinard, L.; Castets, A.; Faure, A.; Lefloch, B. Search for Glycine in the Solar Type Protostar IRAS 16293-2422. *Astron. Astrophys.* **2000**, *362*, 1122–1126.
- (5) Cazaux, S.; Tielens, A. G. G. M.; Ceccarelli, C.; Castets, A.; Wakelam, V.; Caux, E.; Parise, B.; Teyssier, D. The Hot Core around the Low-Mass Protostar IRAS 16293-2422: Scoundrels Rule! *Astrophys. J.* **2003**, *593*, L51–L55.
- (6) Ligterink, N. F. W.; Calcutt, H.; Coutens, A.; Kristensen, L. E.; Bourke, T. L.; Drozdovskaya, M. N.; Müller, H. S. P.; Wampfler, S. F.; van der Wiel, M. H. D.; van Dishoeck, E. F.; et al. The ALMA-PILS Survey: Stringent Limits on Small Amines and Nitrogen-Oxides towards IRAS 16293–2422B. *Astron. Astrophys.* **2018**, *619*, A28 (11 pp).
- (7) Rubin, R. H.; Swenson, G. W., J.; Benson, R. C.; Tigelaar, H. L.; Flygare, W. H. Microwave Detection of Interstellar Formamide. *Astrophys. J.* **1971**, *169*, L39–L44.
- (8) Charnley, S. B.; Tielens, A. G. G. M.; Millar, T. J. On the Molecular Complexity of the Hot Cores in Orion A-Grain Surface Chemistry as “The Last Refuge of the Scoundrel.” *Astrophys. J.* **1992**, *339*, L71–L74.
- (9) Balucani, N.; Ceccarelli, C.; Taquet, V. Formation of Complex Organic Molecules in Cold

- 1
2
3 Objects: The Role of Gas-Phase Reactions. *Mon. Not. R. Astron. Soc.* **2015**, *449*, L16–L20.
4
5
6 (10) Charnley, S. B.; Herbst, E. Reactive Desorption and Radiative Association as Possible
7
8 Drivers of Complex Molecule Formation in the Cold Interstellar Medium. *Astrophys. J.*
9
10 **2013**, *769*, 34 (9 pp).
11
12
13 (11) Garrod, R. T.; Herbst, E. Formation of Methyl Formate and Other Organic Species in the
14
15 Warm-up Phase of Hot Molecular Cores. *Astron. Astrophys.* **2006**, *457*, 927–936.
16
17
18 (12) Öberg, K. I.; Garrod, R. T.; Dishoeck, E. F. van; Linnartz, H. Formation Rates of Complex
19
20 Organics in UV Irradiated CH₃OH-Rich Ices. I. Experiments. *Astron. Astrophys.* **2009**, *504*,
21
22 891–913.
23
24
25 (13) Ruaud, M.; Loison, J. C.; Hickson, K. M.; Gratier, P.; Hersant, F.; Wakelam, V. Modelling
26
27 Complex Organic Molecules in Dense Regions: Eley–Rideal and Complex Induced Reaction.
28
29 *Mon. Not. R. Astron. Soc.* **2015**, *447*, 4004–4017.
30
31
32
33 (14) Watanabe, N.; Kouchi, A. Efficient Formation of Formaldehyde and Methanol by the
34
35 Addition of Hydrogen Atoms to CO in H₂O-CO Ice at 10 K. *Astrophys. J. Lett.* **2002**, *571*,
36
37 L173–L176.
38
39
40 (15) Rimola, A.; Taquet, V.; Ugliengo, P.; Balucani, N.; Ceccarelli, C. Astrophysics Combined
41
42 Quantum Chemical and Modeling Study of CO Hydrogenation on Water Ice. *Astron.*
43
44 *Astrophys.* **2014**, *572*, A70.
45
46
47 (16) Jones, A. P.; Fanciullo, L.; Kohler, M.; Verstraete, L.; Guillet, V.; Bocchio, M.; Ysard, N.
48
49 The Evolution of Amorphous Hydrocarbons in the ISM: Dust Modelling from a New
50
51 Vantage Point. *Astron. Astrophys.* **2014**, *558*, 22.
52
53
54
55 (17) Henning, T. Cosmic Silicates. *Annu. Rev. Astron. Astrophys.* **2010**, *48*, 21–46.
56
57
58 (18) Jones, A. P.; Köhler, M.; Ysard, N.; Bocchio, M.; Verstraete, L. The Global Dust Modelling
59
60 Framework THEMIS. *Astron. Astrophys.* **2017**, *602*, A46 (9 pp).

1
2
3
4
5
6
7
8
9
10
11
12
13
14
15
16
17
18
19
20
21
22
23
24
25
26
27
28
29
30
31
32
33
34
35
36
37
38
39
40
41
42
43
44
45
46
47
48
49
50
51
52
53
54
55
56
57
58
59
60

- (19) Whittet, D. C. B.; Schutte, W. A.; Tielens, A. G. G. M.; Boogert, A. C. A.; de Graauw, T.; Ehrenfreund, P.; Gerakines, P. A.; Helmich, F. P.; Prusti, T.; van Dishoeck, E. F. An ISO SWS View of Interstellar Ices: First Results. *Astron. Astrophys.* **1996**, *360*, L357–L360.
- (20) Boogert, A. C. A.; Gerakines, P. A.; Whittet, D. C. B. Observations of the Icy Universe. *Annu. Rev. Astron. Astrophys.* **2015**, *53*, 541–583.
- (21) Watanabe, N.; Kouchi, A. Ice Surface Reactions: A Key to Chemical Evolution in Space. *Prog. Surf. Sci.* **2008**, *83*, 439–489.
- (22) Fraser, H. J.; Collings, M. P.; Dever, J. W.; McCoustra, M. R. S. Using Laboratory Studies of CO-H₂O Ices to Understand the Non-Detection of a 2152 cm⁻¹ (4.647 μm) Band in the Spectra of Interstellar Ices. *Mon. Not. R. Astron. Soc.* **2004**, *353*, 59–68.
- (23) Zamirri, L.; Casassa, S.; Rimola, A.; Segado-Centellas, M.; Ceccarelli, C.; Ugliengo, P. IR Spectral Fingerprint of Carbon Monoxide in Interstellar Water Ice Models. *Mon. Not. R. Astron. Soc.* **2018**, *480*, 1427–1444.
- (24) McGuire, B. A.; Shingledecker, C. N.; Willis, E. R.; Burkhardt, A. M.; El-Abd, S.; Motiyenko, R. A.; Brogan, C. L.; Hunter, T. R.; Margulès, L.; Guillemin, J.-C.; et al. ALMA Detection of Interstellar Methoxymethanol (CH₃OCH₂OH). *Astrophys. J. Lett.* **2017**, *851*, L46 (8 pp).
- (25) Linnartz, H.; Ioppolo, S.; Fedoseev, G. Atom Addition Reactions in Interstellar Ice Analogues. *Int. Rev. Phys. Chem.* **2015**, *34*, 205–237.
- (26) Barone, V.; Latouche, C.; Skouteris, D.; Vazart, F.; Balucani, N.; Ceccarelli, C.; Lefloch, B. Gas-Phase Formation of the Prebiotic Molecule Formamide: Insights from New Quantum Computations. *Mon. Not. R. Astron. Soc.* **2015**, *453*, L31–L35.
- (27) Vazart, F.; Calderini, D.; Puzzarini, C.; Skouteris, D.; Barone, V. State-of-the-Art Thermochemical and Kinetic Computations for Astrochemical Complex Organic Molecules:

- 1
2
3 Formamide Formation in Cold Interstellar Clouds as a Case Study. *J. Chem. Theory Comput.*
4
5 **2016**, *12*, 5385–5397.
6
7
- 8 (28) Skouteris, D.; Vazart, F.; Ceccarelli, C.; Balucani, N.; Puzzarini, C.; Barone, V. New
9
10 Quantum Chemical Computations of Formamide Deuteration Support Gas-Phase Formation
11
12 of This Prebiotic Molecule. *Mon. Not. R. Astron. Soc.* **2017**, *468*, L1–L5.
13
14
- 15 (29) Wakelam, V.; Loison, J.; Mereau, R.; Ruaud, M. Binding Energies : New Values and Impact
16
17 on the Efficiency of Chemical Desorption. *Mol. Astrophys.* **2017**, *6*, 22–35.
18
19
- 20 (30) Holtom, P. D.; Bennett, C. J.; Osamura, Y.; Mason, N. J.; Kaiser, R. I. A Combined
21
22 Experimental and Theoretical Study on the Formation of the Amino Acid Glycine
23
24 (NH₂CH₂COOH) and Its Isomer (CH₃NHCOOH) in Extraterrestrial Ices. *Astrophys. J.* **2005**,
25
26 *626*, 940–952.
27
28
- 29 (31) Walch, S. P.; Bauschlicher Jr, C. B.; Ricca, A.; Bakes, E. L. O. On the Reaction CH₂O +
30
31 NH₃ → CH₂NH + H₂O. *Chem. Phys. Lett.* **2001**, *333*, 6–11.
32
33
- 34 (32) Redondo, P.; Barrientos, C.; Largo, A. Some Insights into Formamide Formation through
35
36 Gas-Phase Reactions in the Interstellar Medium. *Astrophys. J.* **2013**, *780*, 181 (7 pp).
37
38
- 39 (33) Huang, L. C. L.; Asvany, O.; Chang, A. H. H.; Balucani, N.; Lin, S. H.; Lee, Y. T.; Kaiser,
40
41 R. I. Crossed Beam Reaction of Cyano Radicals with Hydrocarbon Molecules. IV. Chemical
42
43 Dynamics of Cyanoacetylene (HCCCN; X ¹Σ⁺) Formation from Reaction of CN(X ²Σ⁺) with
44
45 Acetylene, C₂H₂(X ¹Σ_g⁺). *J. Chem. Phys.* **2000**, *113*, 8656–8666.
46
47
48
- 49 (34) Navarro-Ruiz, J.; Sodupe, M.; Ugliengo, P.; Rimola, A. Interstellar H Adsorption and H₂
50
51 Formation on the Crystalline (010) Forsterite Surface: A B3LYP-D2* Periodic Study. *Phys.*
52
53 *Chem. Chem. Phys.* **2014**, *16*, 17447–17457.
54
55
- 56 (35) Navarro-Ruiz, J.; Ugliengo, P.; Sodupe, M.; Rimola, A. Does Fe²⁺ in Olivine-Based
57
58 Interstellar Grains Play Any Role in the Formation of H₂? Atomistic Insights from DFT
59
60

1
2
3
4
5
6
7
8
9
10
11
12
13
14
15
16
17
18
19
20
21
22
23
24
25
26
27
28
29
30
31
32
33
34
35
36
37
38
39
40
41
42
43
44
45
46
47
48
49
50
51
52
53
54
55
56
57
58
59
60

Periodic Simulations. *Chem. Commun.* **2016**, *52*, 6873–6876.

- (36) Navarro-Ruiz, J.; Martínez-González, J. Á.; Sodupe, M.; Ugliengo, P.; Rimola, A. Relevance of Silicate Surface Morphology in Interstellar H₂ Formation. Insights from Quantum Chemical Calculations. *Mon. Not. R. Astron. Soc.* **2015**, *453*, 914–924.
- (37) Molpeceres, G.; Rimola, A.; Ceccarelli, C.; Kästner, J.; Ugliengo, P.; Maté, B. Silicate-Mediated Interstellar Water Formation: A Theoretical Study. *Mon. Not. R. Astron. Soc.* **2019**, *482*, 5389–5400.
- (38) Sherrill, C. D. Frontiers in Electronic Structure Theory. *J. Chem. Phys.* **2010**, *132*, 110902 (7 pp).
- (39) Řezáč, J.; Hobza, P. Describing Noncovalent Interactions beyond the Common Approximations: How Accurate Is the “Gold Standard”, CCSD(T) at the Complete Basis Set Limit? *J. Chem. Theory Comput.* **2013**, *9*, 2151–2155.
- (40) Sousa, S. F.; Fernandes, P. A.; Ramos, M. J. General Performance of Density Functionals. *J. Phys. Chem. A* **2007**, *111*, 10439–10452.
- (41) Cramer, C. J.; Truhlar, D. G. Density Functional Theory for Transition Metals and Transition Metal Chemistry. *Phys. Chem. Chem. Phys.* **2009**, *11*, 10757–10816.
- (42) Hao, P.; Sun, J.; Xiao, B.; Ruzsinszky, A.; Csonka, G. I.; Tao, J.; Glindmeyer, S.; Perdew, J. P. Performance of Meta-GGA Functionals on General Main Group Thermochemistry, Kinetics, and Noncovalent Interactions. *J. Chem. Theory Comput.* **2013**, *9*, 355–363.
- (43) Kroes, G.-J. Toward a Database of Chemically Accurate Barrier Heights for Reactions of Molecules with Metal Surfaces. *J. Phys. Chem. Lett.* **2016**, *6*, 4106–4114.
- (44) Grimme, S. Density Functional Theory with London Dispersion Corrections. *WIREs Comput. Mol. Sci.* **2011**, *1*, 211–228.

- 1
2
3 (45) Cramer, C. J. *Essentials of Computational Chemistry*; 2004.
4
5
6 (46) Jensen, F. *Introduction to Computational Chemistry*; 2007.
7
8
9 (47) Atkins, P.; de Paula, J. *Chemical Physics*; 2010.
10
11 (48) Shimonishi, T. Adsorption Energies of Carbon, Nitrogen, and Oxygen Atoms on the Low-
12 Temperature Amorphous Water Ice: A Systematic Estimation from Quantum Chemistry
13 Calculations. *Astrophys. J.* **2018**, *855*, 27 (11 pp).
14
15
16
17
18 (49) Song, L.; Kästner, J. Formation of the Prebiotic Molecule NH₂CHO on Astronomical
19 Amorphous Solid Water Surfaces: Accurate Tunneling Rate Calculations. *Phys. Chem.*
20 *Chem. Phys.* **2016**, *18*, 29278–29285.
21
22
23
24
25 (50) Al-Halabi, A.; Fraser, H. J.; Kroes, G. J.; van Dishoeck, E. F. Adsorption of CO on
26 Amorphous Water-Ice Surfaces. *Astron. Astrophys.* **2004**, *422*, 777–791.
27
28
29
30 (51) Goldman, N.; Reed, E. J.; Fried, L. E.; Kuo, I. W.; Maiti, A. Synthesis of Glycine-Containing
31 Complexes in Impacts of Comets on Early Earth. *Nat. Chem.* **2010**, *2*, 949–954.
32
33
34 (52) Goldman, N.; Tamblyn, I. Prebiotic Chemistry within a Simple Impacting Icy Mixture. *J.*
35 *Phys. Chem. A* **2013**, *117*, 5124–5131.
36
37
38 (53) Laidler, K. J.; King, M. C. Development of Transition-State Theory. *J. Phys. Chem.* **1983**,
39 *87*, 2657–2664.
40
41
42
43 (54) Eyring, H. The Activated Complex in Chemical Reactions. *J. Chem. Phys.* **1935**, *3*, 107–115.
44
45
46 (55) Evans, M. G.; Polanyi, M. Some Applications of the Transition State Method to the
47 Calculation of Reaction Velocities, Especially in Solution. *Trans. Faraday Soc.* **1935**, *31*,
48 875–894.
49
50
51 (56) Williams, D. A. The Interstellar Medium: An Overview. In *Solid State Astrochemistry*;
52 Pirronello, V., Krelowski, J., Manicò, G., Eds.; Proceedings of the NATO Advanced Study
53
54
55
56
57
58
59
60

1
2
3
4
5
6
7
8
9
10
11
12
13
14
15
16
17
18
19
20
21
22
23
24
25
26
27
28
29
30
31
32
33
34
35
36
37
38
39
40
41
42
43
44
45
46
47
48
49
50
51
52
53
54
55
56
57
58
59
60

Institute on Solid State Astrochemistry, 2000; pp 1–20.

- (57) Larson, R. B. *The Evolution of Molecular Clouds*. **1993**.
- (58) Meisner, J.; Kästner, J. Atom Tunneling in Chemistry. *Angew. Chemie Int. Ed.* **2016**, *55*, 5400–5413.
- (59) Eckart, C. The Penetration of a Potential Barrier by Electrons. *Phys. Rev.* **1930**, *35*, 1303–1309.
- (60) Miller, W. H. Semiclassical Limit of Quantum Mechanical Transition State Theory for Nonseparable Systems. *J. Chem. Phys.* **1975**, *62*, 1899–1906.
- (61) Richardson, J. O. Derivation of Instanton Rate Theory from First Principles. *J. Chem. Phys.* **2016**, *144*, 114106 (5 pp).
- (62) Andersson, S.; Nyman, G.; Arnaldsson, A.; Manthe, U.; Jónsson, H. Comparison of Quantum Dynamics and Quantum Transition State Theory Estimates of the H + CH₄ Reaction Rate. *J. Phys. Chem. A* **2009**, *113*, 4468–4478.
- (63) Feynman, R. P. Space-Time Approach to Non-Relativistic Quantum Mechanics. *Rev. Mod. Phys.* **1948**, *20*, 367–387.
- (64) Song, L.; Kästner, J. Formation of the Prebiotic Molecule NH₂CHO on Astronomical Amorphous Solid Water Surfaces: Accurate Tunneling Rate Calculations. *Phys. Chem. Chem. Phys.* **2016**, *18*, 29278–29285.
- (65) Rommel, J. B.; Goumans, T. P. M.; Kästner, J. Locating Instantons in Many Degrees of Freedom. *J. Chem. Theory Comput.* **2011**, *7*, 690–698.
- (66) McQuarrie, D. A.; Simon, J. D. *Physical Chemistry. A Molecular Approach*; University Science Books: Sausalito, CA, USA, 1997.
- (67) Demichelis, R.; Bruno, M.; Massaro, F. R.; Prencipe, M.; de la Pierre, M.; Nestola, F. First-

- 1
2
3 Principle Modelling of Forsterite Surface Properties: Accuracy of Methods and Basis Sets. *J.*
4
5 *Comput. Chem.* **2015**, *36*, 1439–1445.
6
7
- 8 (68) Mukhopadhyay, S.; Bailey, C. L.; Wander, A.; Searle, B. G.; Murny, C. A. Stability of the
9
10 AlF_3 (0 0 -1 2) Surface in H_2O and HF Environments: An Investigation Using Hybrid
11
12 Density Functional Theory and Atomistic Thermodynamics. *Surf. Sci.* **2007**, *601*, 4433–
13
14 4437.
15
16
- 17 (69) Bailey, C. L.; Mukhopadhyay, S.; Wander, A.; Searle, B. G.; Harrison, N. M. Structure and
18
19 Stability of α - AlF_3 Surfaces. *J. Phys. Chem. C* **2009**, *113*, 4976–4983.
20
21
22
- 23 (70) Zamirri, L.; Corno, M.; Rimola, A.; Ugliengo, P. Forsterite Surfaces as Models of Interstellar
24
25 Core Dust Grains: Computational Study of Carbon Monoxide Adsorption. *ACS Earth Sp.*
26
27 *Chem.* **2017**, *1*, 384–398.
28
29
- 30 (71) Chiatti, F.; Corno, M.; Sakhno, Y.; Martra, G.; Ugliengo, P. Revealing Hydroxyapatite
31
32 Nanoparticle Surface Structure by CO Adsorption: A Combined B3LYP and Infrared Study.
33
34 *J. Phys. Chem. C* **2013**, *117*, 25526–25534.
35
36
37
- 38 (72) Boese, A. D.; Sauer, J. Accurate Adsorption Energies for Small Molecules on Oxide
39
40 Surfaces: $\text{CH}_4/\text{MgO}(001)$ and $\text{C}_2\text{H}_6/\text{MgO}(001)$. *J. Comput. Chem.* **2016**, *37*, 2374–2385.
41
42
- 43 (73) Boese, A. D.; Sauer, J. Accurate Adsorption Energies of Small Molecules on Oxide Surfaces:
44
45 $\text{CO-MgO}(001)$. *Phys. Chem. Chem. Phys.* **2013**, *15*, 16481–16493.
46
47
- 48 (74) Escamilla-Roa, E.; Moreno, F. Adsorption of Glycine on Cometary Dust Grains: II — Effect
49
50 of Amorphous Water Ice. *Planet. Space Sci.* **2013**, *75*, 1–10.
51
52
- 53 (75) Escamilla-Roa, E.; Moreno, F. Adsorption of Glycine by Cometary Dust: Astrobiological
54
55 Implications. *Planet. Space Sci.* **2012**, *70*, 1–9.
56
57
- 58 (76) Al-Halabi, A.; Kleyn, A. W.; Van Dishoeck, E. F.; Van Hemert, M. C.; Kroes, G. J. Sticking
59
60 of Hyperthermal CO to the (0001) Face of Crystalline Ice. *J. Phys. Chem. A* **2003**, *107*,

1
2
3 10615–10624.
4

- 5
6 (77) Civalleri, B.; Maschio, L.; Ugliengo, P.; Zicovich-Wilson, C. M. Role of Dispersive
7 Interactions in the CO Adsorption on MgO(001): Periodic B3LYP Calculations Augmented
8 with an Empirical Dispersion Term. *Phys. Chem. Chem. Phys.* **2010**, *12*, 6382–6389.
9
10
11
12 (78) Pisani, C.; Schütz, M.; Casassa, S.; Usvyat, D.; Maschio, L.; Lorenz, M.; Erba, A.
13 CRYSCOR: A Program for the Post-Hartree–Fock Treatment of Periodic Systems. *Phys.*
14 *Chem. Chem. Phys.* **2012**, *14*, 7615–7628.
15
16
17 (79) Chung, L. W.; Sameera, W. M. C.; Ramozzi, R.; Page, A. J.; Hatanaka, M.; Petrova, G. P.;
18 Harris, T. V.; Li, X.; Ke, Z.; Liu, F.; et al. The ONIOM Method and Its Applications. *Chem.*
19 *Rev.* **2015**, *115*, 5678–5796.
20
21
22 (80) Svensson, M.; Humbel, S.; Froese, R. D. J.; Matsubara, T.; Sieber, S.; Morokuma, K.
23 ONIOM: A Multi-Layered Integrated MO + MM Method for Geometry Optimizations and
24 Single Point Energy Predictions. A Test for Diels–Alder Reactions and Pt(P(*t*-Bu)₃)₂+H₂O
25 Oxidative Addition. *J. Phys. Chem.* **1996**, *100*, 19357–19363.
26
27
28 (81) Dapprich, S.; Komáromi, I.; Byun, K. S.; Morokuma, K.; Frisch, M. J. A New ONIOM
29 Implementation in Gaussian 98. 1. The Calculation of Energies, Gradients and Vibrational
30 Frequencies and Electric Field Derivatives. *J. Mol. Struct.* **1999**, *462*, 1–21.
31
32
33 (82) Collings, M. P.; Frankland, V. L.; Lasne, J.; Marchione, D.; Rosu-Finsen, A.; McCoustra, M.
34 R. S. Probing Model Interstellar Grain Surfaces with Small Molecules. *Mon. Not. R. Astron.*
35 *Soc.* **2015**, *449*, 1826–1833.
36
37
38 (83) Collings, M. P.; Anderson, M. A.; Chen, R.; Dever, J. W.; Viti, S.; Williams, D. A.;
39 McCoustra, M. R. S. A Laboratory Survey of the Thermal Desorption of Astrophysically
40 Relevant Molecules. *Mon. Not. R. Astron. Soc.* **2004**, *354*, 1133–1140.
41
42
43 (84) Garrod, R. T. Three-Dimensional, Off-Lattice Monte Carlo Kinetics Simulations of
44
45
46
47
48
49
50
51
52
53
54
55
56
57
58
59
60

- 1
2
3 Interstellar Grain Chemistry and Ice Structure. *Astron. J.* **2013**, *778*, 150 (14 pp).
4
5
6 (85) Sandford, S. A.; Allamandola, L. J.; Tielens, A. G. G. M.; Valero, G. J. Laboratory Studies
7
8 of the Infrared Spectral Properties of CO in Astrophysical Ices. *Astrophys. J.* **1988**, *329*, 498–
9
10 510.
11
12
13 (86) Dendy Sloan Jr., E.; Koh, C. A. *Clathrate Hydrates of Natural Gases*; CRC Press: Boca
14
15 Raton, FL, USA, 2007.
16
17
18 (87) Loveday, J. S.; Nelmes, R. J.; Guthrie, M.; Belmonte, S. A.; Allan, D. R.; Klug, D. D.; Tse, J.
19
20 S.; Handa, Y. P. Stable Methane Hydrate above 2 GPa and the Source of Titan's
21
22 Atmospheric Methane. *Nature* **2001**, *410*, 661–663.
23
24
25 (88) Davidson, D. W.; Desando, M. A.; Cough, S. R.; Handa, Y. P.; Ratcliff, C. I.; Ripmeester, J.
26
27 A.; Tse, J. S. A Clathrate Hydrate of Carbon Monoxide. *Nature* **1987**, *328*, 418–419.
28
29
30 (89) Cossi, M.; Rega, N.; Scalmani, G.; Barone, V. Energies, Structures, and Electronic Properties
31
32 of Molecules in Solution with the C-PCM Solvation Model. *J. Comput. Chem.* **2003**, *24*,
33
34 669–681.
35
36
37 (90) Tomasi, J.; Mennucci, B.; Cammi, R. Quantum Mechanical Continuum Solvation Models.
38
39 *Chem. Rev.* **2005**, *105*, 2999–3094.
40
41
42 (91) Woon, D. E. Ab Initio Quantum Chemical Studies of Reactions in Astrophysical Ices. 4.
43
44 Reactions in Ices Involving HCOOH, CH₂NH, HCN, HNC, NH₃, and H₂O. *Int. J. Quantum*
45
46 *Chem.* **2002**, *88*, 226–235.
47
48
49 (92) Koch, D. M.; Toubin, C.; Peslherbe, G. H.; Hynes, J. T. A Theoretical Study of the
50
51 Formation of the Aminoacetonitrile Precursor of Glycine on Icy Grain Mantles in the
52
53 Interstellar Medium. *J. Phys. Chem. C* **2008**, *112*, 2972–2980.
54
55
56 (93) Riffet, V.; Frison, G.; Bouchoux, G. Quantum-Chemical Modeling of the First Steps of the
57
58 Strecker Synthesis: From the Gas-Phase to Water Solvation. *J. Phys. Chem. A* **2018**, *122*,
59
60

1
2
3
4
5
6
7
8
9
10
11
12
13
14
15
16
17
18
19
20
21
22
23
24
25
26
27
28
29
30
31
32
33
34
35
36
37
38
39
40
41
42
43
44
45
46
47
48
49
50
51
52
53
54
55
56
57
58
59
60

1643–1657.

- (94) Koch, D. M.; Toubin, C.; Xu, S.; Peslherbe, G. H.; Hynes, J. T. Concerted Proton-Transfer Mechanism and Solvation Effects in the HNC/HCN Isomerization on the Surface of Icy Grain Mantles in the Interstellar Medium. *J. Phys. Chem. C* **2007**, *111*, 15026–15033.
- (95) Tielens, A. G. G. M.; Hagen, W. Model Calculations of the Molecular Composition of Interstellar Grain Mantles. *Astron. Astrophys.* **1982**, *114*, 245–260.
- (96) Pirim, C.; Krim, L.; Laffon, C.; Parent, P.; Pauzat, F.; Pilmé, J.; Ellinger, Y. Preliminary Study of the Influence of Environment Conditions on the Successive Hydrogenations of CO. *J. Phys. Chem. A* **2010**, *114*, 3320–3328.
- (97) Woon, D. E. Modeling Gas-Grain Chemistry with Quantum Chemical Cluster Calculations. I. Heterogeneous Hydrogenation of CO and H₂CO on Icy Grain Mantles. *Astrophys. J.* **2002**, *569*, 541–548.
- (98) Rimola, A.; Taquet, V.; Ugliengo, P.; Balucani, N.; Ceccarelli, C. Combined Quantum Chemical and Modeling Study of CO Hydrogenation on Water Ice. *Astron. Astrophys.* **2014**, *572*, A70 (12 pp).
- (99) Goumans, T. P. M.; Wander, A.; Catlow, C. R. A.; Brown, W. A. Silica Grain Catalysis of Methanol Formation. *Mon. Not. R. Astron. Soc.* **2007**, *1832*, 1829–1832.
- (100) Goumans, T. P. M.; Catlow, C. R. A.; Brown, W. A. Hydrogenation of CO on a Silica Surface: An Embedded Cluster Approach. *J. Chem. Phys.* **2008**, *128*, 134709.
- (101) Woon, D. E. Ion-Ice Astrochemistry: Barrierless Low-Energy Deposition Pathways to HCOOH, CH₃OH, and CO₂ on Icy Grain Mantles from Precursor Cations. *Astrophys. J.* **2011**, *728*, 44–49.
- (102) Töpfer, M.; Jusko, P.; Schlemmer, S.; Asvany, O. Double Resonance Rotational Spectroscopy of CH₂D⁺. *Astron. Astrophys.* **2016**, *593*, L11–L14.

- 1
2
3 (103) Roueff, E.; Gerin, M.; Lis, D. C.; Wootten, A.; Marcelino, N.; Cernicharo, J.; Tercero, B.
4
5 CH_2D^+ , the Search for the Holy Grail. *J. Phys. Chem. A* **2013**, *117*, 9959–9967.
6
7
8 (104) Bisschop, S. E.; Jørgensen, J. K.; van Dishoeck, E. F.; de Wachter, E. B. M. Testing Grain-
9
10 Surface Chemistry in Massive Hot-Core Regions. *Astron. Astrophys.* **2007**, *465*, 913–929.
11
12
13 (105) Kahane, C.; Ceccarelli, C.; Faure, A.; Caux, E. Detection of Formamide, the Simplest but
14
15 Crucial Amide, in a Solar-Type Protostar. *Astrophys. J. Lett.* **2013**, *763*, L38 (5 pp).
16
17
18 (106) López-Sepulcre, A.; Jaber, A. A.; Mendoza, E.; Lefloch, B.; Ceccarelli, C.; Vastel, C.;
19
20 Bachiller, R.; Cernicharo, J.; Codella, C.; Kahane, C.; et al. Shedding Light on the Formation
21
22 of the Pre-Biotic Molecule Formamide with ASAI. *Mon. Not. R. Astron. Soc.* **2015**, *449*,
23
24 2438–2458.
25
26
27 (107) Takahiro, Y.; Takano, S.; Watanabe, Y.; Sakai, N.; Sakai, T.; Liu, S.-Y.; Su, Y.-N.; Hirano,
28
29 N.; Takakuwa, S.; Aikawa, Y.; et al. The 3 mm Spectral Line Survey toward the Lynds 1157
30
31 B1 Shocked Region. I. Data. *Publ. Astron. Soc. Japan* **2012**, *64*, 105 (45 pp).
32
33
34 (108) Codella, C.; Ceccarelli, C.; Caselli, P.; Balucani, N.; Barone, V.; Fontani, F.; Lefloch, B.;
35
36 Podio, L.; Viti, S.; Feng, S.; et al. Seeds of Life in Space (SOLIS) II. Formamide in
37
38 Protostellar Shocks: Evidence for Gas-Phase Formation. *Astron. Astrophys.* **2017**, *605*, L3 (7
39
40 pp).
41
42
43 (109) Bianchi, E.; Codella, C.; Ceccarelli, C.; Vazart, F.; Bachiller, R.; Balucani, N.; Bouvier, M.;
44
45 Simone, M. De; Enrique-Romero, J.; Kahane, C.; et al. The Census of Interstellar Complex
46
47 Organic Molecules in the Class I Hot Corino of SVS13-A. *Mon. Not. R. Astron. Soc.* **2019**,
48
49 *483*, 1850–1861.
50
51
52 (110) Bockelée-Morvan, D.; Lis, D. C.; Wink, J. E.; Despois, D.; Crovisier, J.; Bachiller, R.;
53
54 Benford, D. J.; Biver, N.; Colom, P.; Davies, J. K.; et al. New Molecules Found in Comet
55
56 C/1995 O1 (Hale-Bopp) Investigating the Link between Cometary and Interstellar Material.
57
58
59
60

1
2
3
4
5
6
7
8
9
10
11
12
13
14
15
16
17
18
19
20
21
22
23
24
25
26
27
28
29
30
31
32
33
34
35
36
37
38
39
40
41
42
43
44
45
46
47
48
49
50
51
52
53
54
55
56
57
58
59
60

Astron. Astrophys. **2000**, 353, 1101–1114.

- (111) Biver, N.; Bockelée-Morvan, D.; Moreno, R.; Crovisier, J.; Colom, P.; Lis, D. C.; Sandqvist, A.; Boissier, J.; Despois, D.; Milam, S. N. Ethyl Alcohol and Sugar in Comet C/2014 Q2 (Lovejoy). *Sci. Adv.* **2015**, 1, 1–5.
- (112) Spezia, R.; Jeanvoine, Y.; Hase, W. L.; Song, K.; Largo, A. Synthesis of Formamide and Related Organic Species in the Interstellar Medium via Chemical Dynamics Simulations. *Astrophys. J.* **2016**, 826, 107 (8 pp).
- (113) Noble, J. A.; Theule, P.; Congiu, E.; Dulieu, F.; Bonnin, M.; Bassas, A.; Duvernay, F.; Danger, G.; Chiavassa, T. Hydrogenation at Low Temperatures Does Not Always Lead to Saturation: The Case of HNCO. *Astron. Astrophys.* **2015**, 9, A91 (9 pp).
- (114) Rimola, A.; Skouteris, D.; Balucani, N.; Ceccarelli, C.; Enrique-Romero, J.; Taquet, V.; Ugliengo, P. Can Formamide Be Formed on Interstellar Ice? An Atomistic Perspective. *ACS Earth Sp. Chem.* **2018**, 2, 720–734.
- (115) Garrod, R. T.; Weaver, S. L. W.; Herbst, E. Complex Chemistry in Star-Forming Regions: An Expanded Gas-Grain Warm-up Chemical Model. *Astrophys. J.* **2008**, 682, 283–302.
- (116) Öberg, K. I. Photochemistry and Astrochemistry: Photochemical Pathways to Interstellar Complex Organic Molecules. *Chem. Rev.* **2016**, 116, 9631–9663.
- (117) Enrique-Romero, J.; Rimola, A.; Ceccarelli, C.; Balucani, N. The (Impossible ?) Formation of Acetaldehyde on the Grain Surfaces: Insights from Quantum Chemical Calculations. *Mon. Not. R. Astron. Soc.* **2016**, 459, L6–L10.
- (118) Bredehoft, J. H.; Bohler, E.; Schmidt, F.; Borrmann, T.; Swiderek, P. Electron-Induced Synthesis of Formamide in Condensed Mixtures of Carbon Monoxide and Ammonia. *ACS Earth Sp. Chem.* **2017**, 1, 59–59.
- (119) Woon, D. E. Pathways to Glycine and Other Amino Acids in Ultraviolet-Irradiated

- 1
2
3 Astrophysical Ices Determined via Quantum Chemical Modeling. *Astrophys. J.* **2002**, *571*,
4 L177–L180.
5
6
7
8 (120) Rimola, A.; Sodupe, M.; Ugliengo, P. Computational Study of Interstellar Glycine Formation
9 Occurring at Radical Surfaces of Water-Ice Dust Particles. *Astron. J.* **2012**, *754*, 24–33.
10
11
12
13 (121) Park, J.; Woon, D. E. Theoretical Modeling of Formic Acid (HCOOH), FORMATE (HCOO⁻)
14), and Ammonium (NH₄⁺) Vibrational Spectra in Astrophysical Ices. *Astrophys. J.* **2006**, *648*,
15 1285–1290.
16
17
18
19
20 (122) Kayi, H.; Kaiser, R. I.; Head, J. D. A Computational Study on the Structures of
21 Methylamine–carbon Dioxide–water Clusters: Evidence for the Barrier Free Formation.
22 *Phys. Chem. Chem. Phys.* **2011**, *13*, 11083–11098.
23
24
25
26
27 (123) Park, J.; Woon, D. E. Theoretical Investigation of OCN⁻ Charge-Transfer Complexes in
28 Condensed-Phase Media: Spectroscopic Properties in Amorphous Ice. *J. Phys. Chem. A*
29 **2004**, *108*, 6589–6598.
30
31
32
33
34 (124) Park, J.; Woon, D. E. Computational Confirmation of the Carrier for the “XCN” Interstellar
35 Ice Band: OCN⁻ Charge Transfer Complexes. *Astron. J.* **2004**, *601*, L63–L66.
36
37
38
39
40 (125) Snyder, L. E. .; Buhl, D. Observations of Radio Emission from Interstellar Hydrogen
41 Cyanide. *Astrophys. J.* **1971**, *163*, L47–L52.
42
43
44
45 (126) Schilke, P. .; Comito, C. .; Thorwirth, S. First Detection of Vibrationally Excited HNC in
46 Space. *Astrophys. J.* **2003**, *582*, L101–L104.
47
48
49
50 (127) Gardebien, F.; Sevin, A. Catalytic Model Reactions for the HCN Isomerization. I.
51 Theoretical Characterization of Some Water-Catalyzed Mechanisms. *J. Phys. Chem. A* **2003**,
52 *107*, 3925–3934.
53
54
55
56
57 (128) Woon, D. E. A Quantum Chemical Study of the Formation of Cyanide (CN⁻) and Acetate
58 (CH₃COO⁻) Ions in Astrophysical Ices via Proton Transfer from HCN, HNC, or CH₃COOH
59
60

1
2
3
4
5
6
7
8
9
10
11
12
13
14
15
16
17
18
19
20
21
22
23
24
25
26
27
28
29
30
31
32
33
34
35
36
37
38
39
40
41
42
43
44
45
46
47
48
49
50
51
52
53
54
55
56
57
58
59
60

to NH₃. *Comput. Theor. Chem.* **2012**, *984*, 108–112.

- (129) Woon, D. E. Ab Initio Quantum Chemical Studies of Reactions in Astrophysical Ices. 1. Amminolysis, Hydrolysis and Polymerization in H₂CO/NH₃/H₂O Ices. *Icarus* **1999**, *142*, 550–556.
- (130) Woon, D. E. Ab Initio Quantum Chemical Studies of Reactions in Astrophysical Ices 3. Reactions of HOCH₂NH₂ Formed in H₂CO/NH₃/H₂O Ices. *J. Phys. Chem. A* **2001**, *105*, 9478–9481.
- (131) Courmier, D.; Gardebien, F.; Minot, C.; St-Amant, A. A Computational Study of the Water-Catalyzed Formation of NH₂CH₂OH. *Chem. Phys. Lett.* **2005**, *405*, 357–363.
- (132) Rimola, A.; Sodupe, M.; Ugliengo, P. Deep-Space Glycine Formation via Strecker-Type Reactions Activated by Ice Water Dust Mantles. A Computational Approach. *Phys. Chem. Chem. Phys.* **2010**, *12*, 5285–5294.
- (133) Chen, L.; Woon, D. E. A Theoretical Investigation of the Plausibility of Reactions between Ammonia and Carbonyl Species (Formaldehyde, Acetaldehyde, and Acetone) in Interstellar Ice Analogs at Ultracold Temperatures. *J. Phys. Chem. A* **2011**, *115*, 5166–5183.
- (134) Duvernay, F.; Rimola, A.; Theule, P.; Danger, G.; Sanchez, T.; Chiavassa, T. Formaldehyde Chemistry in Cometary Ices: The Case of HOCH₂OH Formation. *Phys. Chem. Chem. Phys.* **2014**, *16*, 24200–24208.
- (135) Fresneau, A.; Danger, G.; Rimola, A.; Duvernay, F.; Theulé, P.; Chiavassa, T. Ice Chemistry of Acetaldehyde Reveals Competitive Reactions in the First Step of the Strecker Synthesis of Alanine Formation of HO–CH(CH₃)–NH₂. *Mon. Not. R. Astron. Soc.* **2015**, *451*, 1649–1660.
- (136) Fresneau, A.; Danger, G.; Rimola, A.; Theulé, P.; Duvernay, F.; Chiavassa, T. Trapping in Water - an Important Prerequisite for Complex Reactivity in Astrophysical Ices: The Case of Acetone (CH₃)₂C=O and Ammonia NH₃. *Mon. Not. R. Astron. Soc.* **2014**, *443*, 2991–3000.

- 1
2
3 (137) Snyder, L. E.; Buhl, D.; Zuckerman, B.; Palmer, P. Microwave Detection of Interstellar
4 Formaldehyde. *Phys. Rev. Lett.* **1969**, *22*, 679–681.
5
6
7
8 (138) Zeng, S.; Quénard, D.; Jiménez-Serra, I.; Martín-Doménech, J. Martín-Pintado, V. M.; L, R.;
9 R, T. First Detection of the Pre-Biotic Molecule Glycolonitrile (HOCH₂CN) in the
10 Interstellar Medium. *Mon. Not. R. Astron. Soc.* **2019**, *484*, L43–L48.
11
12
13
14
15 (139) Woon, D. E. Ab Initio Quantum Chemical Studies of Reactions in Astrophysical Ices. *Icarus*
16 **2001**, *149*, 277–284.
17
18
19
20 (140) Danger, G.; Rimola, A.; Mrad, N. A.; Duvernay, F.; Roussin, G.; Theule, P.; Chiavassa, T.
21 Formation of Hydroxyacetonitrile (HOCH₂CN) and Polyoxymethylene (POM)-Derivatives
22 in Comets from Formaldehyde (CH₂O) and Hydrogen Cyanide (HCN) Activated by Water.
23 *Phys. Chem. Chem. Phys.* **2014**, *16*, 3360–3370.
24
25
26
27
28 (141) Fresneau, A.; Danger, G.; Rimola, A.; Duvernay, F.; Theulé, P.; Chiavassa, T. Thermal
29 Formation of Hydroxynitriles, Precursors of Hydroxyacids in Astrophysical Ice Analogs:
30 Acetone ((CH₃)₂CO) and Hydrogen Cyanide (HCN) Reativity. *Mol. Astrophys.* **2015**, *1*, 1–
31 12.
32
33
34
35 (142) Belloche, A.; Menten, K. M.; Comito, C.; Müller, H. S. P.; Schilke, P.; Ott, J.; Thorwirth, S.;
36 Hieret, C. Detection of Amino Acetonitrile in SgrB2(N). *Astron. Astrophys.* **2008**, *492*, 769–
37 773.
38
39
40 (143) Elsila, J. E.; Glavin, D. P.; Dworkin, J. P. Cometary Glycine Detected in Samples Returned
41 by Stardust. *Meteorit. Planet. Sci.* **2010**, *44*, 1323–1330.
42
43
44
45 (144) Sandford, S. A.; Aléon, J.; Alexander, C. M.; Araki, T.; Bajt, S.; Baratta, G. A.; Borg, J.;
46 Bradley, J. P.; Brownlee, D. E.; Brucato, J. R.; et al. Organics Captured from Comet
47 81P/Wild 2 by the Stardust Spacecraft. *Science* **2006**, *314*, 1720–1724.
48
49
50
51 (145) Altwegg, K.; Balsiger, H.; Bar-Nun, A.; Berthelier, J.-J.; Bieler, A.; Bochsler, P.; Briois, C.;

1
2
3
4
5
6
7
8
9
10
11
12
13
14
15
16
17
18
19
20
21
22
23
24
25
26
27
28
29
30
31
32
33
34
35
36
37
38
39
40
41
42
43
44
45
46
47
48
49
50
51
52
53
54
55
56
57
58
59
60

Calmonte, U.; Combi, M. R.; Cottin, H.; et al. Prebiotic Chemicals – Amino Acid and Phosphorus – in the Coma of Comet 67P/Churyumov-Gerasimenko. *Sci. Adv.* **2016**, *2*, 5 pp.

(146) Pizzarello, S. The Chemistry of Life's Origin: A Carbonaceous Meteorite Perspective. *Acc. Chem. Res.* **2006**, *39*, 231–237.

(147) Strecker, A. Ueber Die Künstliche Bildung Der Milchsäure Und Einen Neuen, Dem Glycocoll Homologen Körper. *Ann. der Chemie und Pharm.* **1850**, *75*, 27–45.

(148) Nhlabatsi, Z. P.; Bhasi, P.; Sitha, S. Possible Interstellar Formation of Glycine from the Reaction of CH₂=NH, CO and H₂O: Catalysis by Extra Water Molecules through the Hydrogen Relay Transport. *Phys. Chem. Chem. Phys.* **2016**, *18*, 375–381.

(149) Nhlabatsi, Z. P.; Bhasi, P.; Sitha, S. Possible Interstellar Formation of Glycine through a Concerted Mechanism: A Computational Study on the Reaction CH₂=NH, CO₂, H₂. *Phys. Chem. Chem. Phys.* **2016**, *18*, 20109–20117.

(150) Lee, H. M.; Choe, J. C. Formation of Glycine from HCN and H₂O: A Computational Mechanistic Study. *Chem. Phys. Lett.* **2017**, *675*, 6–10.

(151) Kayanuma, M.; Kidachi, K.; Shoji, M.; Komatsu, Y.; Sato, A.; Shigeta, Y. A Theoretical Study of the Formation of Glycine via Hydantoin Intermediate in Outer Space Environment. *Chem. Phys. Lett.* **2017**, *687*, 178–183.

(152) Cooper, G. W.; Cronin, J. R. Linear and Cyclic Aliphatic Carboxamides of the Murchison Meteorite: Hydrolyzable Derivatives of Amino Acids and Other Carboxylic Acids. *Geochim. Cosmochim. Acta* **1995**, *59*, 1003–1015.

(153) Shimoyama, A.; Ogasawara, R. Dipeptides and Diketopiperazines in the Yamato-791198 and Murchison Carbonaceous Chondrites. *Orig. life Evol. Biosph.* **2002**, *32*, 165–179.

(154) Vinogradoff, V.; Rimola, A.; Duvernay, F.; Danger, G.; Theulé, P.; Chiavassa, T. The Mechanism of Hexamethylenetetramine (HMT) Formation in the Solid State at Low

- 1
2
3 Temperature. *Phys. Chem. Chem. Phys.* **2012**, *14*, 12309–12320.
4
5
6 (155) Saladino, R.; Crestini, C.; Costanzo, G.; Negri, R.; Di Mauro, E. A Possible Prebiotic
7
8 Synthesis of Purine, Adenine, Cytosine and 4 (3H)-Pyrimidinone from Formamide:
9
10 Implications for the Origin of Life. *Bioorg. Med. Chem.* **2001**, *9*, 1249–1253.
11
12
13 (156) Saladino, R.; Ciambecchini, U.; Crestini, C.; Costanzo, G.; Negri, R.; Di Mauro, E. One-Pot
14
15 TiO₂-Catalyzed Synthesis of Nucleic Bases and Acylonucleosides from Formamide:
16
17 Implications for the Origin of Life. *ChemBioChem* **2003**, *4*, 514–521.
18
19
20 (157) Saladino, R.; Crestini, C.; Costanzo, G.; Di Mauro, E. Advances in the Prebiotic Synthesis of
21
22 Nucleic Acids Bases: Implications for the Origin of Life. *Curr. Org. Chem* **2004**, *8*, 1425–
23
24 1443.
25
26
27 (158) Saladino, R.; Crestini, C.; Ciambecchini, U.; Ciciriello, F.; Costanzo, G.; Di Mauro, E.
28
29 Synthesis and Degradation of Nucleobases and Nucleic Acids by Formamide in the Presence
30
31 of Montmorillonites. *ChemBioChem* **2004**, *5*, 1558–1566.
32
33
34 (159) Saladino, R.; Crestini, C.; Neri, V.; Brucato, J. R.; Colangeli, L.; Ciciriello, F.; Di Mauro, E.;
35
36 Costanzo, G. Synthesis and Degradation of Nucleobases and Nucleic Acids Components by
37
38 Formamide and Cosmic Dust Analogues. *ChemBioChem* **2005**, *6*, 1368–1374.
39
40
41 (160) Saladino, R.; Crestini, C.; Neri, V.; Ciciriello, F.; Costanzo, G.; Di Mauro, E. Origin of
42
43 Informational Polymers: The Concurrent Roles of Formamide and Phosphates.
44
45 *ChemBioChem* **2006**, *7*, 1707–1714.
46
47
48 (161) Saladino, R.; Botta, G.; Bizzarri, B. M.; Di Mauro, E.; Garcia-Ruiz, J. M. A Global Scale
49
50 Scenario for Prebiotic Chemistry: Silica-Based Self-Assembled Mineral Structures and
51
52 Formamide. *Biochemistry* **2016**, *55*, 2806–2811.
53
54
55 (162) Rotelli, L.; Trigo-Rodríguez, J. M.; Moyano-Camero, C. E.; Carota, E.; Botta, L.; Di
56
57 Mauro, E.; Saladino, R. The Key Role of Meteorites in the Formation of Relevant Prebiotic
58
59
60

1
2
3
4
5
6
7
8
9
10
11
12
13
14
15
16
17
18
19
20
21
22
23
24
25
26
27
28
29
30
31
32
33
34
35
36
37
38
39
40
41
42
43
44
45
46
47
48
49
50
51
52
53
54
55
56
57
58
59
60

Molecules in a Formamide/Water Environment. *Sci. Rep.* **2016**, *6*, 38888 (7 pp).

- (163) Nguyen, V. S.; Orlando, T. M.; Leszczynski, J.; Nguyen, M. T. Theoretical Study of the Decomposition of Formamide in the Presence of Water Molecules. *J. Phys. Chem. A* **2013**, *117*, 2543–2555.
- (164) Wang, J.; Gu, J.; Nguyen, M. T.; Springsteen, G.; Leszczynski, J. From Formamide to Purine: A Self-Catalyzed Reaction Pathway Provides a Feasible Mechanism for the Entire Process. *J. Phys. Chem. B* **2013**, *113*, 9333–9342.
- (165) Wang, J.; Gu, J.; Nguyen, M. T.; Springsteen, G.; Leszczynski, J. From Formamide to Purine: An Energetically Viable Mechanistic Reaction Pathway. *J. Phys. Chem. B* **2013**, *117*, 2314–2320.
- (166) Bera, P. P.; Nuevo, M.; Milam, S. N.; Sandford, S. A.; Lee, T. J. Mechanism for the Abiotic Synthesis of Uracil via UV-Induced Oxidation of Pyrimidine in Pure H₂O Ices under Astrophysical Conditions. *J. Chem. Phys.* **2010**, *133*, 104303 (7 pp).
- (167) Bera, P. P.; Nuevo, M.; Materese, C. K.; Sandford, S. A.; Lee, T. J.; Bera, P. P.; Nuevo, M.; Materese, C. K.; Sandford, S. A.; Lee, T. J. Mechanisms for the Formation of Thymine under Astrophysical Conditions and Implications for the Origin of Life. *J. Chem. Phys.* **2016**, *144*, 144308 (7 pp).
- (168) Bera, P. P.; Stein, T.; Head-Gordon, M.; Lee, T. J. Mechanisms of the Formation of Adenine, Guanine, and Their Analogues in UV-Irradiated Mixed NH₃:H₂O Molecular Ices Containing Purine. *Astrobiology* **2017**, *17*, 771–786.
- (169) Burton, S. A.; Stern, J. C.; Elsila, J. E.; Glavin, D. P.; Dworkin, J. P. Understanding Prebiotic Chemistry through the Analysis of Extraterrestrial Amino Acids and Nucleobases in Meteorites. *Chem. Soc. Rev.* **2012**, *41*, 5459–5472.
- (170) Signorile, M.; Salvini, C.; Zamirri, L.; Bonino, F.; Martra, G.; Sodupe, M.; Ugliengo, P.

1
2
3 Formamide Adsorption at the Amorphous Silica Surface: A Combined Experimental and
4
5 Computational Approach. *Life* **2018**, 8, 42 (13 pp).
6
7
8
9
10
11
12
13

14 **For TOC only**
15
16

